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## THE SCUBA-2 COSMOLOGY LEGACY SURVEY: MULTI-WAVELENGTH PROPERTIES OF ALMA-IDENTIFIED SUBMILLIMETER GALAXIES IN UKIDSS-UDS

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### ABSTRACT

We present a multi-wavelength analysis of 52 sub-millimeter galaxies (SMGs), identified using ALMA 870  $\mu\text{m}$  continuum imaging in a pilot program to precisely locate bright SCUBA2-selected sub-mm sources in the UKIDSS Ultra Deep Survey (UDS) field. Using the available deep (especially near-infrared), panoramic imaging of the UDS field at optical-to-radio wavelengths we characterize key properties of the SMG population. The median photometric redshift of the bright ALMA / SCUBA-2 UDS (AS2UDS) SMGs that are detected in a sufficient number of wavebands to derive a robust photometric redshift is  $z = 2.65 \pm 0.13$ . However, similar to previous studies, 27% of the SMGs are too faint at optical-to-near-infrared wavelengths to derive a reliable photometric redshift. Assuming that these SMGs lie at  $z \gtrsim 3$  raises the median redshift of the full sample to  $z = 2.9 \pm 0.2$ . A subset of 23, unlensed, bright AS2UDS SMGs have sizes measured from resolved imaging of their rest-frame far-infrared emission. We show that the extent and luminosity of the far-infrared emission are consistent with the dust emission arising from regions that are, on average, optically thick at a wavelength of  $\lambda_0 \geq 75 \mu\text{m}$  ( $1-\sigma$  dispersion of 55–90  $\mu\text{m}$ ). Using the dust masses derived from our optically-thick spectral energy distribution models we determine that these galaxies have a median hydrogen column density of  $N_H = 9.8^{+1.4}_{-0.7} \times 10^{23} \text{ cm}^{-2}$ , or a corresponding median  $V$ -band obscuration of  $A_V = 540^{+80}_{-40}$  mag, averaged along the line of sight to the source of their restframe  $\sim 200 \mu\text{m}$  emission. We discuss the implications of this extreme attenuation by dust for the multiwavelength study of dusty starbursts and reddening-sensitive tracers of star formation.

*Subject headings:* galaxies: starburst, galaxies: high redshift

### 1. INTRODUCTION

In the local Universe, Ultra Luminous Infrared Galaxies (ULIRGs), with far-infrared luminosities of  $\geq 10^{12} L_\odot$ , represent the most intense sites of ongoing star formation (e.g. Soifer et al. 1984). Despite ULIRGs

having immense star formation rates of  $\gtrsim 100 M_\odot \text{ yr}^{-1}$  the bolometric luminosity of these galaxies is dominated by emission from inter-stellar dust grains, which obscure the ongoing starburst at ultra-violet and optical wavelengths and re-emit in the far-infrared. The dust emission takes the form of a modified blackbody function, which at the typical temperatures of these sources ( $T_d \sim 40 \text{ K}$ ; e.g. Symeonidis et al. 2013) peaks at 50–100  $\mu\text{m}$  restframe and declines strongly at longer wavelengths. The strong increase in flux density with decrease wavelength results in the well-known negative  $k$ -correction at sub-millimeter (sub-mm) wavelengths; for sources at higher redshift the observed sub-mm waveband samples the spectral energy distribution (SED) closer to its peak. Indeed, the negative  $k$ -correction is so strong at sub-mm wavelengths ( $\sim 850 \mu\text{m}$ ) that a galaxy with a constant luminosity and temperature has an almost constant brightness with redshift, as the increase in the flux density of the source directly counters the effect of cosmological dimming out to  $z \sim 7$  (see Blain et al. 2002). Observations at sub-mm wavelengths thus provide a unique tracer of obscured star formation across a large fraction of the age of the Universe.

The first deep, extragalactic surveys at sub-mm wavelengths, undertaken with bolometer cameras on single-dish facilities, unveiled a population of bright sources at flux densities of  $S_{850} \gtrsim 5\text{--}15 \text{ mJy}$  (e.g. Smail et al. 1997; Barger et al. 1998; Hughes et al. 1998; Eales et al. 1999; Greve et al. 2004; Coppin et al. 2006; Scott et al. 2008). While these surveys detected only a modest number of

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sub-mm sources the surface density of these detections was used to infer that the number of far-infrared-bright galaxies must undergo a rapid evolution with redshift (e.g. Smail et al. 1997). However, the low angular resolution of single-dish facilities (typically  $\sim 15''$ ) means that identifying the sub-mm galaxies (SMGs;  $S_{850} > 1 \text{ mJy}$ ) that are responsible for each sub-mm source is not possible without significant assumptions about the properties of these sources at other wavelengths. Typically, the correlation between radio and far-infrared emission (e.g. Ivison et al. 1998, 2000) was exploited to provide statistical identification of the counterparts to sub-mm sources, since facilities such as the Very Large Array (VLA) can provide the arcsecond resolution imaging that is required to identify individual galaxies (Ivison et al. 2002, 2004, 2007; Bertoldi et al. 2007; Biggs et al. 2011; Lindner et al. 2011).

Identifying single-dish-detected sub-mm sources through observations with the VLA at 1.4 GHz paved the way for our understanding of the SMG population, with the initial analysis of these radio-identified SMGs confirming the high-redshift nature (median  $z \sim 2.5$ ; Chapman et al. 2005) and ULIRG-like far-infrared luminosities ( $\geq 10^{12} L_{\odot}$ ; e.g. Magnelli et al. 2012) of the SMG population. Further analysis has shown that SMGs are relatively massive, gas-rich galaxies (gas masses of  $\sim 5 \times 10^{10} M_{\odot}$ , e.g. Bothwell et al. 2013) with space densities of  $\sim 10^{-5} \text{ Mpc}^{-3}$  (Hainline et al. 2011), while rest-frame optical imaging from *HST* has demonstrated that the visible stellar component in SMGs has a disturbed or irregular morphology (e.g. Conselice et al. 2003; Chapman et al. 2003; Swinbank et al. 2010a; Targett et al. 2013; Wiklind et al. 2014; Chen et al. 2014). Thus, SMGs appear to have similar properties to local ULIRGs, despite being  $\sim 10^3$  times more numerous than these proposed analogs, at a fixed far-infrared luminosity (e.g. Chapman et al. 2005; Lindner et al. 2011; Magnelli et al. 2012; Yun et al. 2012; Swinbank et al. 2014).

Although identifying sub-mm sources at radio wavelengths has proven to be a powerful tool to understand the SMGs population the technique is susceptible to issues with mis-identification and incompleteness, problems that are inherent in any analysis involving statistical associations. Thus, recent interferometric observations undertaken at sub-mm / mm wavelengths with facilities such as the Sub-mm Array (SMA) and Plateau de Bure Interferometer (PdBI), or, more recently, with the Atacama Large sub-/Millimeter Interferometer (ALMA) have improved our understanding of the SMG population (e.g. Younger et al. 2009; Wang et al. 2011; Smolčić et al. 2012; Hodge et al. 2013; Karim et al. 2013; Walter et al. 2016; Dunlop et al. 2017). Crucially, these facilities can provide imaging at sub-mm wavelengths with an angular resolution of only a few arcseconds, or better, thus providing the sub-arcsecond positional accuracy that is required to identify the multi-wavelength counterparts to single-dish identified sub-mm sources and circumventing the requirement for statistical associations at other wavelengths. Furthermore, observations with ALMA can achieve sufficiently high resolution to resolve SMGs at sub-mm wavelengths, enabling the internal processes that govern the obscured starburst to be studied and

allowing a direct comparison with local ULIRGs.

In the first, large interferometric study of sub-mm sources undertaken, Hodge et al. (2013) used ALMA to obtain sensitive, high resolution images of 126 sub-mm sources that were identified in the LABOCA survey of the Extended *Chandra* Deep Field-South (LESS), identifying 99 SMGs within the  $17.3''$ -diameter primary beam of 88 of the highest quality ALMA observations. These observations confirmed previous suggestions that a significant fraction of bright single-dish-identified sub-mm sources are comprised of a blend of multiple individual SMGs (e.g. Wang et al. 2011) and led to the suggestion that the intrinsic  $870 \mu\text{m}$  number counts may have a strong decline at  $> 9 \text{ mJy}$ , potentially indicating a maximal luminosity to high redshift starbursts (Karim et al. 2013; but see also Barger et al. 2012; Chen et al. 2013). However, while ALMA-LESS (ALESS) represents a complete sample of sub-mm sources ( $S_{870\mu\text{m}} > 4.4 \text{ mJy}$ ) the steep shape of the single-dish  $870 \mu\text{m}$  number counts means that only 10 sources brighter than  $S_{870\mu\text{m}} 9 \text{ mJy}$  were observed as part of the survey.

To investigate the properties of the brightest, unlensed SMGs we undertook a pilot ALMA survey of 30 bright sub-mm sources (Simpson et al. 2015b) that were identified as part of the SCUBA-2 Cosmology Legacy Survey (Geach et al. 2017). These 30 sources are located in the Ultra Deep Survey (UDS) field, the deepest component of the panoramic UKIRT Infrared Deep Sky Survey (UKIDSS; Lawrence et al. 2007), and thus have deep multi-wavelength imaging across optical-to-near-infrared wavelengths. In previous work, we presented the source catalog, number counts and far-infrared morphologies of the 52 SMGs that were detected in these 30 ALMA maps (see Simpson et al. 2015a,b). We demonstrated that  $61^{+19}_{-15} \%$  of single-dish-identified sub-mm are comprised of two or more SMGs ( $> 1 \text{ mJy}$ ) and that the number density of these secondary sources is inconsistent with them being chance line-of-sight projections (Simpson et al. 2015b). Furthermore, we used our high-resolution ALMA imaging to show that the far-infrared region in SMGs has a median angular size (deconvolved FWHM of the major axis) of  $0.30 \pm 0.04''$  (Simpson et al. 2015a).

Here, we use the available photometric imaging of the UDS field to study the properties of these 52 ALMA-identified SMGs at optical to radio wavelengths, including an analysis of the dust properties of the 23 SMGs that were resolved in our  $870 \mu\text{m}$  ALMA imaging. In particular, the sample of SMGs studied here doubles the number of bright  $850 \mu\text{m}$  sources that have been interferometrically-identified using ALMA and we use the improved statistics that this provides to search for trends in the SMG population with flux density. The paper is structured as follows. In §2 we discuss our sample selection. In §3 we describe the multi-wavelength coverage of our ALMA-identified SMGs and our SED fitting procedures, before discussing the multi-wavelength properties of these SMGs in §4. In §5 we present the redshift distribution and far-infrared properties of the AS2UDS SMGs. Furthermore, we discuss the dust properties of the 23 SMGs with measured sizes at observed  $870 \mu\text{m}$  and present the implications for the optical depth and attenuation of stellar light in these sources. We

discuss these in §5 and give our main conclusions in §6. We adopt a cosmology with  $H_0 = 67.8 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_\Lambda = 0.69$ , and  $\Omega_m = 0.31$  (*Planck* Collaboration et al. 2014). Throughout this work error estimates are from a bootstrap analysis and all magnitudes are in the AB system (Oke 1974), unless otherwise stated.

## 2. SAMPLE SELECTION

In this work we study the multi-wavelength properties of a sample of 52 SMGs that were identified using targeted ALMA Band 7 continuum imaging of 30 bright single-dish-detected sub-mm sources. Here we give a brief overview of the sample selection from the initial single-dish imaging and the ALMA data reduction. For a detailed description of the data reduction and analysis we refer the reader to Simpson et al. (2015b). The initial sample of 30 sub-mm sources was detected in wide-field SCUBA-2  $850 \mu\text{m}$  imaging of the UDS field, taken as part of the SCUBA-2 Cosmology Legacy Survey (S2CLS; Geach et al. 2017). Our sample was constructed for the ALMA Cycle-1 deadline in early 2013 from a preliminary version of this S2CLS map, which reached a  $1-\sigma$  depth of  $2.0 \text{ mJy}$ . From this early map we selected 30 apparently bright sources detected at  $> 4\sigma$  for ALMA follow-up observations. Crucially, the ALMA follow-up observations targeted each sub-mm source at the same wavelength and provide imaging across a primary beam that encompasses the SCUBA-2 beam ( $\text{FWHM} = 14.8''$ ), but with a synthesized beam that is a factor of  $\gtrsim 400$  smaller.

All 30 SCUBA-2-detected sub-mm sources in our sample were observed with ALMA on 2013 November 1 with 26 12 m antennae. The array configuration yielded a synthesized beam (using Briggs weighting with robust = 0.5) of  $0.35'' \times 0.25''$  and the data were calibrated and imaged using the COMMON ASTRONOMY SOFTWARE APPLICATION (CASA; version 4.2.1). We note that two versions of the final, cleaned maps were produced: a “high resolution” set of maps with a median  $1-\sigma$  depth of  $0.21 \text{ mJy beam}^{-1}$  and a median synthesized beam of  $0.35'' \times 0.25''$ ; and a set of  $uv$ -tapered “detection” images with a median  $1-\sigma$  depth of  $0.26 \text{ mJy beam}^{-1}$  and a median synthesized beam of  $0.80'' \times 0.65''$ . Simpson et al. (2015b) construct a source catalog of 52 SMGs ( $S_{870} = 1.3\text{--}12.9 \text{ mJy}$ ) that are detected at  $> 4\sigma$  in the 30 ALMA “detection” images. A subset of 23 / 52 SMGs are detected at a sufficiently high SNR ( $> 10$ ) in the “high resolution” images to allow a study of their morphology at observed  $870 \mu\text{m}$ .

## 3. OBSERVATIONS

The focus of this paper is a multi-wavelength analysis of the 52 SMGs detected at  $\text{SNR} > 4$  in our deep ALMA imaging as part of a pilot study for a large survey of  $\sim 10^3$  SMGs with ALMA in the S2CLS UDS map (Stach et al. in prep.). Crucially, this pilot AS2UDS sample is comprised of 17 SMGs with  $870 \mu\text{m}$  flux densities brighter than  $7.5 \text{ mJy}$ , a factor of two increase relative to previous ALMA surveys of  $870 \mu\text{m}$  sources (ALESS; Hodge et al. 2013), and we use this increase in dynamic range to study both bright sources and the overall SMG population ( $S_{850} \gtrsim 1 \text{ mJy}$ ). The UKIDSS UDS is a target field for deep, panoramic observations from optical-to-radio wavelengths and we use the existing archival

images from these multi-wavelength surveys throughout our analysis. In the following we give a description of each of these surveys and how we use the resulting data products to analyze our sample of SMGs.

### 3.1. Optical and Near-Infrared Imaging

The dust enshrouded nature of SMGs means that deep near-infrared imaging is essential for determining properties such as their photometric redshifts. The UKIDSS observations of the  $\sim 0.8 \text{ deg}^2$  UDS comprise four Wide-Field Camera (WFCAM) pointings, in the  $J$ -,  $H$ -, and  $K$ -bands. In this paper we use the images and catalogs released as part of the UKIDSS data release 8 (DR8). The DR8 release contains data taken between 2005 and 2010, and the final  $J$ -,  $H$ -, and  $K$ -band mosaics have a median  $5-\sigma$  depth ( $2''$  apertures) of  $J = 24.9$ ,  $H = 24.2$  and  $K = 24.6$ , respectively.

Deep observations of the UDS have also been taken in the  $U$ -band with Megacam at the Canada France-Hawaii Telescope (CFHT) and in the  $B$ ,  $V$ ,  $R$ ,  $i'$ , and  $z'$ -bands with Suprime-cam at the Subaru telescope. The Megacam / CFHT  $U$ -band imaging reaches a  $5\sigma$  ( $2''$  diameter aperture) depth of  $U = 26.75$  (Foucaud et al. in prep.) and the Suprime-cam imaging has a limiting  $3\sigma$  depth of  $B = 28.4$ ,  $V = 27.8$ ,  $R = 27.7$ ,  $i' = 27.7$ ,  $z' = 26.7$  in the  $B$ ,  $V$ ,  $R$ ,  $i'$ , and  $z'$ -bands, respectively ( $2''$  diameter apertures; Furusawa et al. 2008). Furthermore, deep *Spitzer* data, obtained as part of the SpUDS program (PI: J. Dunlop), provides imaging reaching a  $5\sigma$  depth of  $m_{3.6} = 24.2$  and  $m_{4.5} = 24.0$  at  $3.6 \mu\text{m}$  and  $4.5 \mu\text{m}$ , respectively.

The DR8 UKIDSS catalog contains the  $U$ -to- $4.5 \mu\text{m}$  photometry for  $\sim 140,000$  sources detected in the deep  $K$ -band image of the UDS. For each source, eleven band photometry was determined by running SExtractor (Bertin & Arnouts 1996) in “dual-image” mode on the images described above, using the UKIDSS  $K$ -band image as the detection image. The flux of each source was measured in a  $3''$ -diameter aperture and to ensure consistent galaxy colors, aperture corrections that account for source-blending were applied to  $U$ -band,  $3.6 \mu\text{m}$  and  $4.5 \mu\text{m}$  photometry. Hartley et al. (2013) use the color-matched photometry to derive photometric redshifts for the sources in the UKIDSS UDS catalog but, to allow a direct comparison to previous studies, we apply a further correction to convert the  $3''$  aperture flux measurements to a “total” magnitude. We stack 15 isolated stars in the  $K$ -band image and determine a  $3''$ -to-total aperture correction of  $-0.2 \text{ mag}$ , which we apply to the UKIDSS photometry.

#### 3.1.1. Photometric Redshifts

Photometric redshifts were determined for the sources in the UKIDSS UDS DR8 catalog, using the eleven-band optical-to-near-infrared photometry described in §3.1. The analysis was presented in Hartley et al. (2013) and Mortlock et al. (2013), but we give a summary here. The template fitting code EAZY (Brammer et al. 2008) was used to fit a library of seven template spectral energy distributions (SEDs) to the photometry of each  $K$ -band-selected source in the DR8 release. First, a subset of 2146 spectroscopically-confirmed sources (excluding AGN) were used to calibrate the photometric redshifts,

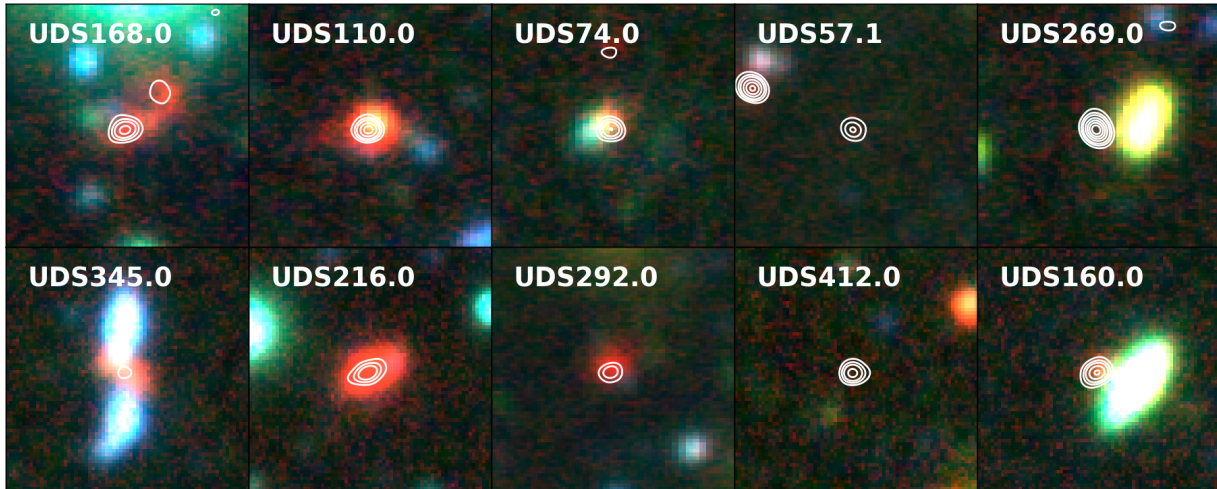


FIG. 1.— Example  $10'' \times 10''$  true color images (constructed from  $B$ ,  $I$  and  $K$ ) of ten SMGs in our ALMA-identified sample. The sources are selected to be representative of the optical-to-near-infrared properties of the full sample and from left-to-right the columns show sources with disturbed / irregular morphologies, compact sources and optically-blank sources. The final column shows two of the four SMGs in our sample that we identify as gravitationally lensed sources. The SMGs typically appear red in color, although we note that  $27 \pm 7\%$  of the sample are not detected in the deep UKIDSS  $K$ -band imaging ( $5\text{-}\sigma$  depth  $K = 24.6$ ).

and correct for any zero-point offsets between the template SEDs and the UDS photometry. The majority of these spectroscopic redshifts are drawn from the ESO large program UDSz (Almaini et al. in prep targeting  $z > 1$  galaxies), but a small number of archival redshifts are included. The redshift of each spectroscopically-confirmed source was fixed at the spectroscopic redshift in the SED fitting and the offsets between the template and observed fluxes were used to iteratively correct the zero-points of each of the eleven filters. An offset of 0.15 mag was applied to the  $U$ -band photometry and the offsets in all remaining bands were  $\leq 0.05$  mag.

The final photometric redshifts for the spectroscopic sample are found to have a median  $(z_{\text{phot}} - z_{\text{spec}}) / (1 + z_{\text{spec}}) = 0.020$ , with a  $1\text{-}\sigma$  dispersion of 0.031, indicating very good agreement between the redshifts (catastrophic outliers at  $(z_{\text{phot}} - z_{\text{spec}}) / (1 + z_{\text{spec}}) > 0.15$  were removed). We note that as AS2UDS SMGs represent a distinct population of highly dust obscured galaxies the accuracy of photometric redshifts for these sources may be lower than estimated for the overall UKIDSS catalog. Indeed, previous studies have shown that for SMGs with comparable photometry the  $1\text{-}\sigma$  dispersion in  $(z_{\text{phot}} - z_{\text{spec}}) / (1 + z_{\text{spec}})$  is typically  $\lesssim 0.1$  (Simpson et al. 2014; Danielson et al. 2016). Crucially, these studies do not find any bias in the photometric redshifts, and have demonstrated good agreement between the photometric and spectroscopic redshifts of SMGs. Further details regarding analysis and reliability testing of the UKIDSS photometric redshift catalog are given in Hartley et al. (2013) and Mortlock et al. (2013).

### 3.2. Far-infrared Imaging

The UDS field was observed at 250, 350, and  $500 \mu\text{m}$  with the Spectral and Photometric Imaging Receiver (SPIRE) onboard the *Herschel Space Observatory* as part of the *Herschel* Multi-tiered Extragalactic Survey (HerMES; Oliver et al. 2012). Observations of the field were taken in seven “sub-blocks”, each with an exposure time of 5.4 ks, resulting in a total exposure time for the field of 37.8 ks. As described in Swinbank et al. (2014),

we retrieved the Level 2 data products from *Herschel* European Space Agency archive and aligned and co-added the images. To ensure the co-added SPIRE images are aligned with the astrometric reference frame of the deep radio imaging of the UDS (see § 3.3) we stacked the maps at the VLA radio positions, centroided the stacked emission and applied shifts of  $< 1.5''$  to each SPIRE map.

The SPIRE / *Herschel* imaging has an angular resolution of  $\sim 18, 25, \text{ and } 35''$  at 250, 350, and  $500 \mu\text{m}$ . The coarse resolution of the imaging means that it is vital that we consider the effect of source blending when determining the far-infrared flux densities of the SMGs in our sample. To determine accurate flux densities for the SMGs we deblend the UDS maps following the procedure presented in Swinbank et al. (2014), which includes extensive tests to confirm the reliability and completeness of the analysis. First, we use the available  $24 \mu\text{m}$  / *Spitzer* source catalogs ( $> 5\sigma$ ) to construct a catalog of likely infrared-bright galaxies that are used as positional priors in the deblending. The positions for 52 SMGs from ALMA are added to the prior list and any duplicate sources within  $1.5''$  are removed from the final prior catalog, ensuring that the ALMA positions are retained. The SPIRE maps are then deblended by fitting the appropriate beam at the position of each source in the prior catalog and minimizing the  $\chi^2$  statistic. To ensure that they do not “over-deblend” the longer-wavelength, coarser resolution SPIRE imaging, Swinbank et al. (2014) deblend the maps in order of increasing wavelength and only include  $24 \mu\text{m}$  sources that are detected at  $> 2\sigma$  in the shorter wavebands as positional priors. Upper limits for non-detections and appropriate error bars are determined through simulations (see Swinbank et al. 2014). The detection fractions of the AS2UDS SMGs are 24 / 48, 26 / 48, 19 / 48, at 250, 350, and  $500 \mu\text{m}$ , respectively (25 / 48 detected in  $\geq 2$  wavebands), and the deblended SPIRE fluxes and associated uncertainties are given in Table 1.

#### 3.2.1. Far-infrared SED Fitting

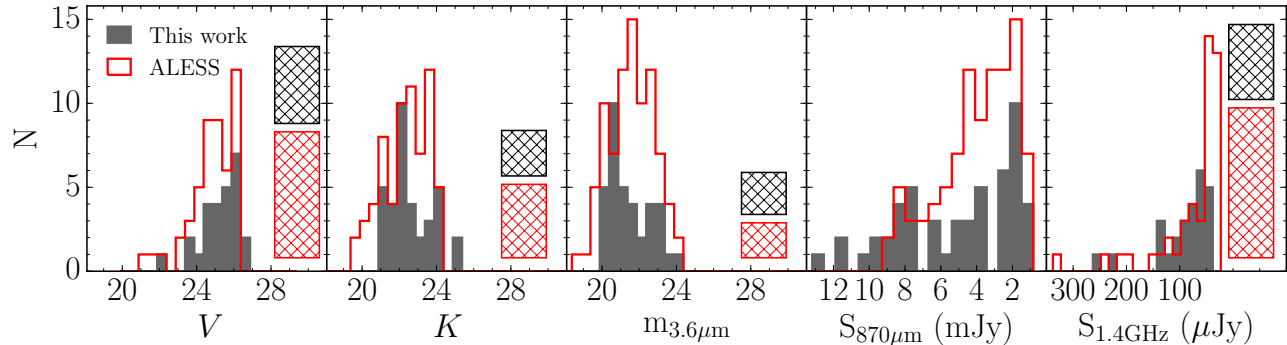


FIG. 2.— The apparent magnitude distributions of the AS2UDS sample of SMGs in the  $V$ ,  $K$  and  $3.6 \mu\text{m}$  wavebands, along with their flux density distributions at  $870 \mu\text{m}$  and  $1.4 \text{GHz}$ . The median  $V$ ,  $K$  and IRAC  $3.6 \mu\text{m}$  apparent magnitudes, including the numbers of non-detections (hatched regions), are  $V = 26.4^{+\infty}_{-0.3}$ ,  $K = 23.0^{+0.7}_{-0.5}$  and  $m_{3.6} = 21.8^{+0.6}_{-0.3}$ . For comparison we show the magnitude distributions of the ALMA-identified SMGs in the ALESS sample (Simpson et al. 2014). The ALESS SMGs have a median  $S_{870} = 3.5 \pm 0.3 \text{ mJy}$  and so are marginally fainter than the SMGs in our sample, median  $S_{870} = 4.2^{+0.9}_{-0.6} \text{ mJy}$ . The  $1.4 \text{GHz}$  VLA imaging of the UDS reaches a  $1-\sigma$  depth of  $\sim 7 \mu\text{Jy beam}^{-1}$ , at its deepest, and in total 27/52 SMGs from AS2UDS are detected with a median flux density of  $S_{1.4\text{GHz}} = 42^{+11}_{-42} \mu\text{Jy}$  (Arumugam et al. in prep.). The SMGs in our sample are marginally brighter at  $1.4 \text{GHz}$  than the ALESS SMGs (median  $S_{1.4\text{GHz}} < 19.5 \mu\text{Jy}$ ), which we attribute to the differences in the  $870 \mu\text{m}$  flux density distribution of the two samples.

To characterize the temperatures and far-infrared luminosities ( $8\text{--}1000 \mu\text{m}$ ) of the SMGs in our sample we fit the observed far-infrared photometry of each source with a single-temperature, modified blackbody function

$$S_{\nu_{\text{obs}}} \propto (1 - e^{-\tau_{\nu_{\text{rest}}}}) \times B(\nu_{\text{rest}}, T) \quad (1)$$

where  $B(\nu, T)$  represents the Planck function,  $\tau_{\nu} = (\frac{\nu}{\nu_0})^{\beta}$  is the frequency dependent optical depth of the dust,  $\nu_0$  is the frequency at which the optical depth is unity,  $z$  is the redshift of the source, and  $\beta$  is the dust emissivity. In our analysis we adopt  $\beta = 1.8$ , consistent with previous studies of the far-infrared emission from SMGs, and in-line with studies of galactic dust emission presented by the *Planck* collaboration (Planck Collaboration et al. 2011).

The optical depth and the dust temperature parameters in the modified blackbody function are correlated; both a decrease in the optical-depth, and an increase in the dust temperature, shifts the peak of the SED bluewards. To allow a direct comparison to previous work we first make the assumption that the dust emission originates from regions that are optically-thin (i.e.  $\nu_0 \gg \nu$ ), simplifying the modified blackbody function to

$$S_{\nu_{\text{obs}}} \propto \nu_{\text{rest}}^{\beta} \times B(\nu_{\text{rest}}, T). \quad (2)$$

However, as we discuss in § 5.2 the emission from SMGs does not originate from regions that are optically-thin, which is consistent with studies of far-infrared-bright sources in the local Universe (e.g. Scoville et al. 2017). As such, the dust temperature derived using the optically-thin approximation does not represent the true temperature of the dust emission regions and in the following work we refer to it as a characteristic dust temperature. We first compare the characteristic dust temperatures of the AS2UDS SMGs to samples of local sources that have been analyzed in the same manner before estimating the true dust temperatures of these SMGs in § 5.2.

We fit the optically-thin modified blackbody function (Equation 2) to the photometry of each SMG in our sample that has a photometric redshift, using an affine-

invariant, Markov Chain Monte Carlo (MCMC) sampler (EMCEE; Foreman-Mackey et al. 2013). By using an MCMC approach to model the far-infrared emission we can include the full redshift probability distribution function for each SMG, and thus determine robust uncertainties for each model parameter. The MCMC code is run using 50 “walkers” for a total  $10^6$  steps, following an initial, and conservative, burn-in phase of  $10^4$  steps. An analysis of the time-series data indicates that for each source the burn-in phase is complete and the chain is well-mixed. As discussed in § 4.1, a number of the SMGs in our sample are not detected in some, or all, of the SPIRE wavebands. To account for non-detections in the SED fitting we adopt the modification to the  $\chi^2$  statistic presented by Sawicki (2012)

$$\chi_{\text{mod}}^2 = \sum_i \left( \frac{f_{d,i} - f_{m,i}}{\sigma_i} \right)^2 - 2 \sum_j \ln \left\{ \sqrt{\frac{\pi}{2}} \sigma_j \left[ 1 + \text{erf} \left( \frac{f_{lim,j} - f_{m,j}}{\sqrt{2} \sigma_j} \right) \right] \right\} \quad (3)$$

where the summations over  $i$  and  $j$  represent wavebands in which a source is detected or non-detected, respectively;  $f_d$  is the observed flux density of a source;  $f_m$  is the model flux density;  $\sigma$  is the uncertainty on the measured flux density; and  $f_{lim}$  is the upper limit on the flux density of the source in the relevant waveband. If a source is detected in all wavebands then the summation over  $j$  vanishes and the statistic reverts to the standard  $\chi^2$  statistic. However, if a source is not detected in the  $j$ th waveband then the modification to  $\chi^2$  includes the probability that the source would be considered a non-detection in the imaging, given the current value of the model. If an SMG is not detected in any of the SPIRE wavebands then we can only determine a plausible range for its far-infrared luminosity, which is determined by the maximum characteristic dust temperature that is consistent with the SPIRE upper limits and the temperature of the Cosmic Microwave Background (CMB) at the photometric redshift of the source. To calculate this

range we fix the SED at the measured  $870\mu\text{m}$  flux density and determine the minimum and maximum parameter values that produce a model in agreement with all upper limits.

The SED model contains three parameters; the normalization,  $N$ ; the characteristic dust temperature  $T_d$ ; and the redshift of the source,  $z$ . The well-known degeneracy between temperature and redshift means that we cannot constrain both parameters without prior information (Blain & Longair 1996). Thus, we use the full redshift probability distribution for each source, as determined in the optical-to-near-infrared SED fitting, as a prior on the redshift. We note that we place an additional, flat prior on the characteristic dust temperature of each source that ensures that it is higher than the lower limit set by the temperature of the CMB, at the appropriate redshift. Finally, it is well-known a single-temperature modified blackbody function fails to reproduce short wavelength ( $\lesssim 50\mu\text{m}$ ) dust emission from an infrared-bright galaxy, where emission from increasingly warm dust results in a power law flux distribution (Blain et al. 2002). We caution that we do not account for this in our analysis and that a single temperature modified blackbody typically under-estimates the total far-infrared luminosity of a source by 20%, relative to empirical galaxy template SEDs (e.g. Swinbank et al. 2014).

### 3.3. Radio / 1.4 GHz Imaging

The UDS field was observed by the VLA at 1.4 GHz as part of the project UDS20 (Arumugam et al. in prep.). A total of 14 pointings were used to mosaic an area of  $\sim 1.3\text{deg}^2$  centered on the UDS field. The observations were taken in A, B, and C–D configuration, yielding a typical synthesized beam of  $\sim 1.8''$  FWHM. The final map reaches a  $1-\sigma$  depth of  $7\mu\text{Jy beam}^{-1}$  at its deepest and  $\sim 7000$  sources are detected across the field at a peak  $\text{SNR} > 4$ .

We match our ALMA catalog to the 1.4 GHz catalog and identify 26 matches within  $2''$  (maximum separation  $0.9''$ ; expected false matching rate  $< 0.1\%$ ). However, two bright SMGs (UDS156.0 and 156.1;  $S_{870} = 8.5$  and  $9.7\text{mJy}$ , respectively) have a small on-sky separation of  $2.3''$ . We inspect the VLA imaging at the position of these sources and identify a bright 1.4 GHz source that is centered directly between the position of both SMGs and extended in the direction of both sources. We estimate the flux density of each SMG by fitting two Gaussian profiles centered at the positions of the ALMA sources.

Hence, in total 27/52 ALMA-identified SMGs are detected in the deep 1.4 GHz imaging with flux densities ranging from  $40\text{--}780\mu\text{Jy}$  (Table 1). The median flux density of the sample is weakly constrained at  $42^{+11}_{-42}\mu\text{Jy}$ . We note that the astrometry of the ALMA and VLA images is well-aligned, with median offsets between the ALMA and VLA source positions of  $-0.08^{+0.03}_{-0.02}$  in R. A. and  $-0.03^{+0.05}_{-0.03}$  in Decl.

## 4. ANALYSIS

We first study the fundamental observable characteristics of our SMG sample before determining their redshifts, which allow us to determine key physical properties such as the epoch of their activity. An extensive

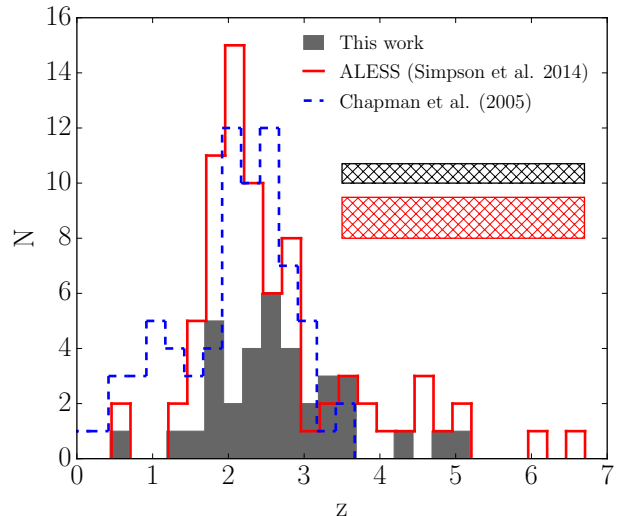


Fig. 3.— The photometric redshift distribution of the ALMA-identified SMGs in our sample. The 35 SMGs in AS2UDS that have sufficient photometry to derive a photometric redshift have a median redshift of  $z = 2.65 \pm 0.13$ . For comparison we show the photometric redshift distribution of ALMA-identified SMGs in the ECDF-S (ALESS; Simpson et al. 2014) and the spectroscopic redshift distribution of radio-identified SMGs presented by Chapman et al. (2005). We find that the median redshift of the SMGs in our sample is marginally higher than for the ALESS SMGs,  $z = 2.31^{+0.08}_{-0.13}$ . However, the median values are consistent at the  $1.5-\sigma$  confidence level, and the shape of the distributions appear to be in agreement. Similarly, the radio-identified sample presented in Chapman et al. (2005) lie at a lower median redshift of  $z = 2.20 \pm 0.10$ , and have notably more sources at  $z < 1$ . Hatched regions represent the 13 and 19 SMGs in the AS2UDS and ALESS samples, respectively, that have insufficient photometry to derive a reliable photometric redshift.

literature search reveals that none of the SMGs in our sample have archival spectroscopic redshifts (including from UDSz; § 3.1.1). However, we can make use of the excellent multi-wavelength imaging that is available in the UDS (see § 3.1) and the photometric redshift estimates that have been derived from the UKIDSS UDS (Hartley et al. 2013). In the following section we present the multi-wavelength properties of our sample of AS2UDS SMGs and compare these to other samples of ALMA-identified SMGs.

### 4.1. Optical and Near-infrared Photometry

To determine the optical-to-near-infrared photometry of the SMGs in our sample we match the ALMA-identified positions to the UKIDSS  $K$ -band catalog. A matching radius of  $1''$  was adopted (85% of matches are found within  $\lesssim 0.5''$ ) to account for both the formal uncertainty on the ALMA positions ( $\sigma \sim 0.14''$  for a  $4\sigma$  detection; see Ivison et al. 2007) and any intrinsic spatial offset resulting from dust obscuration ( $\sigma \sim 0.3''$ , with offsets of up to  $2''$  to individual components; Chen et al. 2015). To ensure that the ALMA and UKIDSS astrometric reference frames are well-aligned, we compare the positions of the 33 matched sources in both catalogs. We identify a small astrometric offset between the reference frames of  $0.09^{+0.05}_{-0.04}$  and  $-0.15^{+0.04}_{-0.07}$  in R.A. and Decl., which we apply to the UKIDSS UDS astrometry. Note that we then repeated the source matching using the astrometrically aligned catalog, but did not identify any further matches to the sources in our sample at  $< 1''$ .

To ensure that we have not missed any potential counterparts to the AS2UDS SMGs we extend the search radius for counterparts to  $2''$ ; consistent with previous high resolution studies of SMGs that have demonstrated significant positional offsets between the observed  $870\ \mu\text{m}$  and near-infrared emission in a fraction of counterparts as a result of the high dust obscuration, disturbed morphology and often structured dust regions that are typical of the SMG population (e.g. Chen et al. 2015). Matching the ALMA and UKIDSS catalogs, we identify a potential counterpart to both UDS 199.1 and UDS 269.1, at separations of  $1.3''$  and  $1.6''$ , respectively. To test the reliability of these proposed counterparts we first construct a catalog of 50,000 random positions within the area of the UKIDSS  $K$ -band image. We match our fake source list to the UKIDSS source catalog and estimate a false-matching-rate of 8% and 12% at  $1.3''$  and  $1.6''$ , respectively. However, previous studies of the redshift distribution of SMGs have indicated that the majority of sources lie at  $z > 1.5$  and we can use this prior knowledge in our analysis. Thus, we repeat our analysis and estimate that at a separation of  $1.3''$  and  $1.6''$  the false-matching-rate of a source in our catalog of random positions to a  $z > 1.5$  source in the UKIDSS catalog is 3% and 4%, respectively. Both of the proposed counterparts to both UDS 199.1 and UDS 269.1 lie at  $z > 1.5$  (see Table 2) thus, given the low likelihood that these are spurious matches, we include both in our analysis.

In Figure 1 we show example  $BIK$  true-color images for ten SMGs that span the full range of ALMA  $870\ \mu\text{m}$  flux density for our catalog. The images demonstrate that if an SMG is detected in the optical-to-near-infrared imaging it typically appears red in the  $BIK$  color images. The observed  $V$ ,  $K$ , and  $3.6\ \mu\text{m}$  magnitude distributions of the SMGs in our sample are shown in Figure 2.

It is important to note that the counterparts to the SMGs are identified by matching to a  $K$ -band-selected catalog. The depth of the  $K$ -band image relative to the IRAC imaging ( $5\sigma$  depths of  $K = 24.6$  and  $m_{3.6} = 24.2$  mag) means that we do not expect to have missed a significant number of additional counterparts in the longer wavelength imaging, except for the very reddest sources. Indeed, we examine the IRAC imaging and only identify counterparts to a further four SMGs at  $3.6\ \mu\text{m}$  and / or  $4.5\ \mu\text{m}$  (UDS 57.1, UDS 199.0, UDS 286.2, and UDS 412.0). However, these sources are not detected at any other wavelengths and, as shown in Simpson et al. (2014), detections in at least four wave-bands are required to determine even crude photometric redshifts; a crucial first step towards understanding the physical properties of these sources. We note that three of these SMGs are not detected in the available  $1.4\ \text{GHz}$  imaging, and that the far-infrared emission from all four SMGs appears to peak redwards of  $350\ \mu\text{m}$ , indicating that they likely lie at higher redshift ( $z \gtrsim 3$ ; see Swinbank et al. 2014).

As our ALMA observations targeted bright sub-mm sources ( $S_{850} \gtrsim 8\ \text{mJy}$ ) we must be aware of the influence of gravitational lensing on our initial selection (e.g. Blain 1996; Chapman et al. 2002). To quantify the effect of gravitational lensing on our sample we visually inspected the optical imaging of all 52 SMGs, identifying four sources (UDS 109.0, 160.0, 269.0 & 286.0) as being

potentially gravitationally lensed. All four of these SMGs lie close to, but are spatially offset from, galaxies at  $z < 1$  (see Figure 1). These SMGs are faint, or undetected, at optical wavelengths, relative to the foreground sources, and although the emission in the IRAC imaging typically appears extended from the bright galaxy in the direction of the SMG it is heavily blended. None of the SMGs show evidence of being multiply imaged indicating that the potential magnification factors are likely to be modest. We highlight these four sources in Table 1 and do not include them in our main analysis.

The median apparent magnitudes of the sample are  $V = 26.4^{+\infty}_{-0.3}$ ,  $K = 23.0^{+0.7}_{-0.5}$ , and  $m_{3.6} = 21.8^{+0.6}_{-0.3}$ . Excluding gravitationally lensed sources,  $27 \pm 7\%$  of the sample (13 / 48 SMGs) are undetected in the deep UKIDSS UDS imaging ( $K \leq 24.6$  mag). As expected for dusty, high redshift sources the counterpart detection rate decreases in bluer wavebands, falling to  $54 \pm 8\%$  (26 / 48) in the  $B$ -band. For comparison, in Figure 2 we show the magnitude distributions for the 96 ALESS SMGs (Simpson et al. 2014). The ALESS SMGs (Hodge et al. 2013) were identified in ALMA  $870\ \mu\text{m}$  follow-up imaging of single-dish-identified  $870\ \mu\text{m}$  sub-mm sources and are well-matched to the sample presented here. The parent sample for the AS2UDS SMGs is brighter at  $870\ \mu\text{m}$  than the ALESS SMGs, and this is reflected in the  $870\ \mu\text{m}$  flux densities of the sources (median  $S_{870} = 4.2^{+0.9}_{-0.6}$  mJy and  $S_{870} = 3.5 \pm 0.3$  mJy for AS2UDS and ALESS, respectively). The ALESS SMGs have median apparent magnitudes of  $V = 26.1^{+0.2}_{-0.1}$ ,  $K = 23.0^{+0.3}_{-0.4}$ , and  $m_{3.6} = 21.8^{+0.2}_{-0.1}$ , respectively, in good agreement with the observed magnitude distributions of the AS2UDS SMGs.

#### 4.2. Optically-faint SMGs

We next investigate whether the detectability of counterparts to SMGs in the  $K$ -band is a function of  $870\ \mu\text{m}$  flux density. The  $K$ -band detected sources in our sample have a median  $S_{870} = 4.2^{+1.0}_{-0.6}$  mJy, compared to a median  $S_{870} = 2.3^{+1.9}_{-0.2}$  mJy for the non-detections; a small hint, albeit statistically insignificant, that the  $K$ -band non-detections may be fainter at  $870\ \mu\text{m}$ . To investigate this further we combine the AS2UDS and ALESS samples and repeat the analysis, but to ensure a fair comparison we consider any AS2UDS SMGs fainter than detection limit of the  $K$ -band imaging of the ALESS SMGs ( $K \leq 24.4$ ) as non-detected. The median  $870\ \mu\text{m}$  flux densities for the  $K$ -band detections and non-detections in the combined sample are  $S_{870} = 4.0 \pm 0.3$  mJy and  $S_{870} = 2.3^{+0.3}_{-0.2}$  mJy, respectively, again suggesting that the  $K$ -band undetected SMGs are fainter at  $870\ \mu\text{m}$  at the  $2.8\sigma$  significance-level. If this result is confirmed in larger samples then these fainter SMGs represent either the lower-luminosity (either due to higher dust obscuration or lower stellar mass) and / or high redshift tail to the SMG population. As discussed by Simpson et al. (2014) placing these SMGs at low redshift introduces a strong bi-modality into the distribution of rest-frame  $H$ -band luminosity (a proxy for stellar mass) or dust obscuration in the SMG population. This problem can be avoided by instead assuming that these sources simply represent the high-redshift tail to the SMG population that lie below the detection threshold of the optical-to-near-infrared imaging. Hence, in § 5.1 we discuss the impact of placing



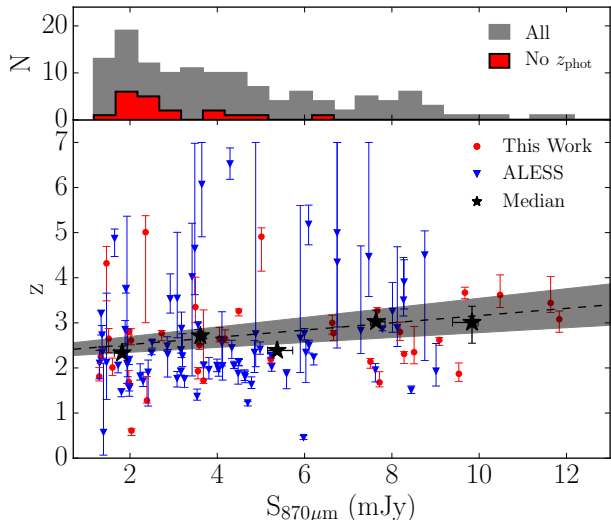


FIG. 4.— The photometric redshifts of the 35 SMGs presented in this work as a function of their  $870\ \mu\text{m}$  flux densities. For comparison we also show the 77 ALESS SMGs with photometric redshifts detected in ALMA imaging of single-dish sources in the ECDF-S (Hodge et al. 2013; Simpson et al. 2014). We combine both samples of SMGs and plot the median of the combined sample in  $2\ \text{mJy}$ -wide bins. A trend of increasing flux density with increasing redshift is observed for the SMGs with photometric redshift estimates, and indeed a linear fit to the data shows a slope of  $0.080 \pm 0.026$  (dashed line and shaded region represent the best-fit and 68% confidence region respectively). However, in the upper panel we show the flux density distribution for the SMGs that do not have a photometric redshift estimates and the overall sample. If we assume that these optically-faint SMGs lie at  $z > 3.0$  (a likely hypothesis; Simpson et al. 2014) then the observed trend in  $870\ \mu\text{m}$  flux density with redshift weakens, yielding a best-fit slope  $0.000 \pm 0.001$ , and thus is consistent with no evolution with cosmic time. We therefore conclude that there is currently no evidence for a trend of redshift with  $870\ \mu\text{m}$  flux density for SMGs.

these SMGs at high redshift.

## 5. RESULTS & DISCUSSION

### 5.1. Photometric Redshift Distribution of SMGs

The 35 SMGs from our sample of 48 that are detected in the  $K$ -band imaging of the UDS have a median redshift of  $z_{\text{phot}} = 2.65 \pm 0.13$ . The shape of the redshift distribution is slightly skewed to high redshift and extends to  $z \sim 5$  (Figure 3)<sup>14</sup>. We first compare the redshift distribution of the AS2UDS SMGs to a sample of radio-identified sub-mm sources with spectroscopic redshifts presented by Chapman et al. (2005). The Chapman et al. (2005) sample of SMGs lie at a median redshift of  $z = 2.20 \pm 0.10$ , slightly lower than the redshift of the SMGs presented here. An offset between the redshift distribution of the radio-selected and  $870\ \mu\text{m}$ -selected SMGs is expected due to the respective positive and negative  $k$ -corrections in each wave-band. To ensure a fair comparison, we consider the 21 / 35 SMGs in our red-

<sup>14</sup> A number of the AS2UDS SMGs have large redshift uncertainties, or secondary minima in their redshift probability distribution functions. To investigate whether the overall redshift distribution is sensitive to these we create a single redshift probability distribution for the sample by co-adding the integral-normalized redshift probability function of each SMG. The shape of the combined redshift probability function is well-matched to the shape of the redshift distribution shown in the Figure 3, and corresponds to a median redshift of  $z_{\text{phot}} = 2.61^{+0.07}_{-0.13}$ , in agreement with the median redshift of the AS2UDS SMGs.

shift distribution that are detected in the VLA 1.4 GHz imaging presented here, which we note has a comparable depth to radio imaging employed by Chapman et al. (2005;  $7\ \mu\text{Jy beam}^{-1}$  here versus  $\sim 9\ \mu\text{Jy beam}^{-1}$ ). These radio-detected, ALMA-identified SMGs have a median redshift of  $z_{\text{phot}} = 2.62^{+0.15}_{-0.31}$ , slightly higher than the sample presented by Chapman et al. (2005), but consistent at the  $1\text{-}\sigma$  confidence level. We note that the median redshift of the radio-identified subset of the AS2UDS SMGs sample is consistent with the  $K$ -band detected subset, indicating that for the SMGs presented here the radio selection limit is well-matched to the depth of the  $K$ -band image.

Next, we use the photometric redshifts that we have determined for the AS2UDS SMGs to test whether multiple sources that are detected in the same ALMA map tend to lie at the same photometric redshift. Thus testing if these SMGs are physically-associated or are simply line-of-sight projections. Due to the large associated uncertainties on the photometric redshift of any individual SMG (median  $\Delta z \sim 0.4$ ) we cannot test whether the SMGs located in the same map are physically associated on a source-by-source basis. Instead, we sample the full redshift probability distribution for each SMG and search for statistical over-densities of sources at the same redshift in each ALMA map, relative to the overall population. We find that the AS2UDS SMGs that are detected in the same ALMA map are  $17 \pm 9\%$  more likely to lie at  $\Delta z < 0.4$ , compared to SMGs that are detected in a different ALMA map. While this provides tentative evidence that a fraction of these SMGs are physically-associated we caution that this is a  $2\sigma$  results, and that the test can only be performed for the eleven pairs where photometric redshifts are available for both SMGs.

In Figure 3 we compare the redshift distribution of the AS2UDS SMGs to the photometric redshift distribution of the 77 ALESS SMGs presented by Simpson et al. (2014). The ALESS SMGs lie at a median redshift of  $z = 2.3 \pm 0.1$  and we note that the shape of the distribution is similar to the results presented here; there is a dearth of SMGs in both samples at  $z \lesssim 1$ , and a high redshift tail extends to  $z > 3$ . A further 19 ALESS SMGs are detected in an insufficient number of optical-to-near-infrared wave-bands to determine a photometric redshift. The fraction of SMGs in our sample without photometric redshift estimates is  $27^{+10}_{-7}\%$  (13 / 48), which is consistent with that for the ALESS sample ( $20 \pm 5\%$ ) at the  $< 1\text{-}\sigma$  confidence level, assuming Poisson statistics.

The median redshift of the SMGs presented in this work is marginally higher than the ALESS SMGs. The key difference between the samples is that the AS2UDS SMGs are brighter, on average, at  $870\ \mu\text{m}$  than the ALESS sample, and have a significantly higher fraction of more luminous sources ( $29^{+8}_{-7}\%$  at  $S_{870} > 7.5\ \text{mJy}$ , compared to  $9^{+4}_{-3}\%$  for the ALESS SMGs; see Figure 2). Thus, a possible explanation for the higher median redshift of the SMGs presented here, relative to ALESS, is that brighter SMGs are preferentially found at higher redshift. Indeed, a number of authors have previously suggested that  $870\text{-}\mu\text{m}$ -brighter sources may lie at higher redshift (e.g. Ivison et al. 2002, 2007; Koprowski et al. 2014).

To investigate whether there is evidence for such a trend we combine the AS2UDS and ALESS SMGs, and analyze the combined sample of 144 sources. As shown in Figure 4 we find the SMGs *that have photometric redshift estimates* do exhibit a positive trend of increasing flux density with redshift and a linear fit to the data finds a slope of  $0.080 \pm 0.026$ . However, we strongly caution that this trend is mirrored by a decrease in the redshift completeness with decreasing  $870\text{-}\mu\text{m}$  flux; 22/32 of the SMGs that do not have a photometric redshift have  $S_{870} < 3\text{ mJy}$ .

As discussed previously, the optically-faint SMGs that do not have a photometric redshift estimate are likely to lie at higher redshifts than the average AS2UDS SMGs. So, if these SMGs are conservatively placed at  $z=3\text{--}6$  then the positive trend between  $S_{870}$  and redshift is no longer apparent, and a linear fit to the data returns a slope of  $-0.000 \pm 0.001$ . Placing these optically-faint SMGs at  $z > 3.0$  does however raise the median redshifts of the AS2UDS and ALESS samples to  $z = 2.9 \pm 0.2$  and  $z = 2.5 \pm 0.2$ , respectively. As such the median redshift of the AS2UDS SMG is  $\Delta z \sim 0.4$  higher than that found for the ALESS SMGs, when non-detections are treated in the same manner. As discussed above, this disparity in the redshift distribution of these samples of SMGs is not due to a difference in the flux density distribution of both samples. Instead, it probably indicates that there is a difference in the underlying distribution of galaxies in the ECDF-S and UDS fields, reinforcing the conclusion that the redshift distribution of SMGs is sensitive to the large-scale-structure of the Universe (Williams et al. 2011).

### 5.2. Far-infrared Properties

As described in §3.2.1 we estimate the far-infrared luminosities and characteristic dust temperatures of the AS2UDS SMGs by fitting an optically-thin modified blackbody to the observed photometry of each source. In total, 24 AS2UDS SMGs are detected at a sufficient number of optical-to-far-infrared wavelengths that we can estimate both their far-infrared luminosities and characteristic dust temperatures (i.e. detected in at least one SPIRE waveband and have a photometric redshift) and these SMGs have a median far-infrared luminosity and characteristic dust temperature of  $L_{\text{FIR}} = 3.4_{-0.2}^{+0.2} \times 10^{12} L_{\odot}$  and  $T_d = 32.9_{-1.6}^{+0.4}\text{ K}$ , respectively. In Table 2 we provided the range of plausible far-infrared luminosities for the 11 AS2UDS SMGs that are not detected in the SPIRE imaging, but have photometric redshift estimates. We compare the far-infrared luminosities and characteristic dust temperatures of the sources presented here to the ALESS sample of SMGs but, to ensure an accurate comparison, we repeat the far-infrared SED fitting for these ALESS SMGs using the photometry presented by Swinbank et al. (2014) and the SED fitting method presented here. The 59 ALESS SMGs that are detected in at least one SPIRE waveband have a median far-infrared luminosity and characteristic dust temperature of  $L_{\text{FIR}} = 3.2_{-0.7}^{+0.3} \times 10^{12} L_{\odot}$  and  $T_d = 32.1_{-1.0}^{+1.3}\text{ K}$ , respectively, consistent with the median properties of the AS2UDS SMGs. The similarity between the median luminosities and the characteristic dust temperatures of the AS2UDS and the ALESS SMGs is unsurprising, given the close agreement between the median

$870\text{-}\mu\text{m}$  flux density of the samples and the relevant depth of the SPIRE multi-wavelength coverage, but confirms the homogeneity in the properties of bright,  $870\text{-}\mu\text{m}$ -selected sources.

In Figure 5, we investigate the relationship between the luminosity and characteristic dust temperature of the AS2UDS SMGs and find a clear trend of increasing luminosity with temperature. However, before comparing this result to other samples we must consider the selection function of our  $870\text{-}\mu\text{m}$  ALMA observations. Thus, in Figure 5 we also show the far-infrared selection function at the depth of our  $870\text{-}\mu\text{m}$  observations, as a function of both characteristic dust temperature and redshift. We see the well-known negative  $k$ -correction in the sub-mm waveband that results in near-uniform selection in far-infrared luminosity out to  $z \sim 6$ , for sources at a fixed characteristic dust temperature of  $\sim 30\text{ K}$ . However, this uniform selection with redshift does not hold for all dust temperatures. Crucially, at the median redshift of our sample our observations are  $\sim 3 \times$  less sensitive to sources that have a characteristic dust temperature of  $T_d = 40\text{ K}$ , relative to sources with  $T_d = 30\text{ K}$  (see also Blain et al. 2002). As such care must be taken when comparing the results presented here to samples selected at a different observed wavelength and / or redshift.

It has been suggested that SMGs may be the high redshift analogues of ULIRGs that are seen in the local Universe. To investigate whether low-redshift populations can be used as templates for SMGs, we compare the far-infrared properties of the AS2UDS SMGs to the SPIRE-selected sample of  $z \sim 0\text{--}1$  U / LIRGs presented by Symeonidis et al. (2013). We convert the far-infrared luminosities from Symeonidis et al. (2013) to account for the difference between modeling the emission with empirical templates and a modified blackbody, but note that the characteristic dust temperatures are derived in the same manner as the present work. The local sample shows a clear trend of decreasing peak SED wavelength, or equally increasing characteristic dust temperature, with far-infrared luminosity, and the median characteristic dust temperature of the sample rises from  $T_d = 29\text{--}39\text{ K}$  over the range  $L_{\text{FIR}} = 0.1\text{--}2.5 \times 10^{12} L_{\odot}$ . The AS2UDS SMGs exhibit a similar trend between far-infrared luminosity and characteristic dust temperature, relative to the local sample, but appear to be significantly cooler at a fixed luminosity; the AS2UDS SMGs with far-infrared luminosities of  $L_{\text{FIR}} = 3\text{--}5 \times 10^{12} L_{\odot}$  have a median dust temperature of  $T_d = 32 \pm 2\text{ K}$ .

The selection function for the observations of local sources is such that the sample is effectively complete in characteristic dust temperature. As discussed above, this is not true for the AS2UDS SMGs as the  $850\text{-}\mu\text{m}$  selection results in a bias towards cooler sources. Thus, a direct comparison between the low- and high-redshift samples over a wide-range in luminosity is not possible (see Figure 5). Nevertheless, we can consider the AS2UDS SMG with far-infrared luminosities of  $L_{\text{FIR}} = 3\text{--}5 \times 10^{12} L_{\odot}$ , where, based on their temperature distribution we would expect to detect sources with the same range in  $T_d$  as those that are seen in the local Universe. At these luminosities the AS2UDS SMGs have a median characteristic dust temperature of  $T_d = 32 \pm 2\text{ K}$ , and are thus  $\gtrsim 7\text{ K}$  cooler than the most luminous subset of the local sample,

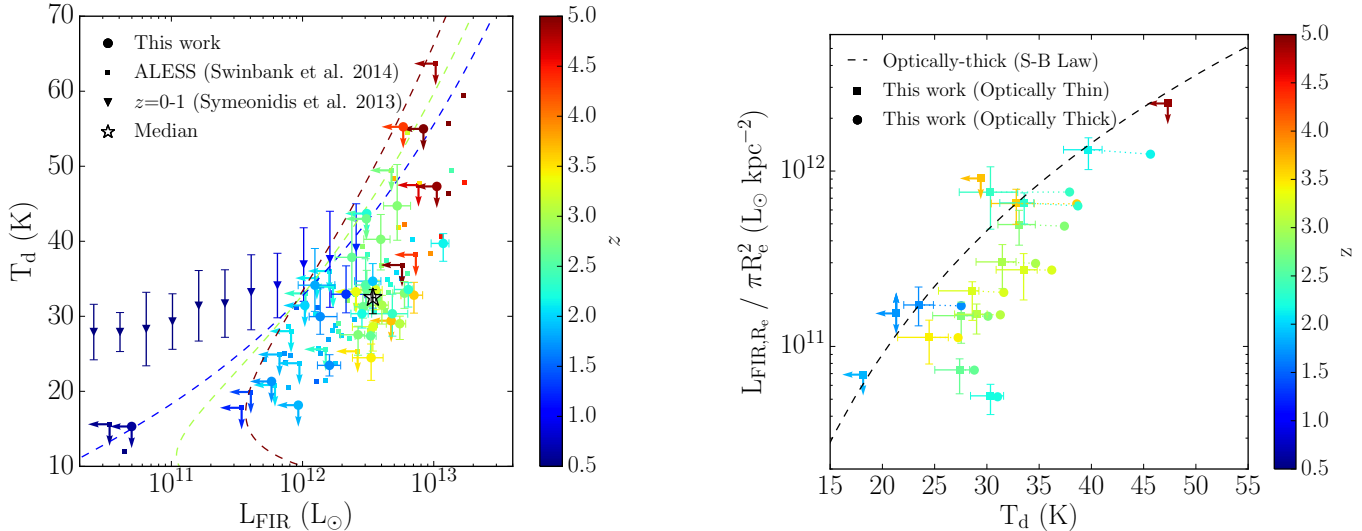


FIG. 5.— *Left*: The characteristic dust temperature of the AS2UDS SMGs as a function of far-infrared luminosity, and color-coded by photometric redshift. For comparison we show the ALESS sample (Swinbank et al. 2014) and a sample of infrared-bright sources at  $z < 1$ , selected from observations with *Herschel*/SPIRE (Symeonidis et al. 2013). The AS2UDS show a clear trend of increasing characteristic dust temperature with increasing luminosity, consistent with observations of low redshift LIRGs/ULIRGs, but appear cooler at a fixed luminosity. We caution that due to our selection at  $870\ \mu\text{m}$ , and the depth of our ALMA observations, the samples do not overlap significantly in far-infrared luminosity; the selection function of our observations at  $870\ \mu\text{m}$  ( $S_{870} \geq 1.4\ \text{mJy}$ ), as a function of redshift, is represented by dashed lines for  $z = 1, 3, 5$ . Nevertheless, the AS2UDS SMGs appear to have characteristic dust temperatures that are  $\sim 8\ \text{K}$  cooler than sources with comparable far-infrared luminosities at  $z \lesssim 1$ . The lower characteristic dust temperatures of the AS2UDS SMGs, at a fixed luminosity, suggest that the dust emission from these sources arises from regions with significantly different physical properties to low redshift far-infrared-bright galaxies, emphasizing that these sources cannot be simply used as analogues to describe high redshift SMGs. *Right*: The luminosity density of the 18 AS2UDS SMGs that have both photometric redshifts and measured sizes from high resolution  $870\ \mu\text{m}$  imaging with ALMA. The data show a clear trend of increasing luminosity density with dust temperature, consistent with the Stefan-Boltzmann law for blackbody emission (dashed line) if the dust regions in the sources are assumed to be optically-thin. However, the Stefan-Boltzmann law is only valid if the far-infrared emission originates from dust clouds that are optically-thick at all far-infrared wavelengths. We show that the AS2UDS SMGs are optically-thick at a median  $\lambda_0 > 75\ \mu\text{m}$ , with a  $1\text{-}\sigma$  dispersion of  $55\text{--}90\ \mu\text{m}$ , resulting in a systematic increase in the implied intrinsic dust temperatures of each source by  $\geq 3.1^{+1.0}_{-0.3}\ \text{K}$ , on average.

or  $\geq 8\ \text{K}$  if we extrapolate the best-fit relation to the local sample to match the median far-infrared luminosity of the AS2UDS SMGs. Including the ALESS SMGs in this comparison raises the median characteristic dust temperature of the SMGs to  $T_d = 33 \pm 1\ \text{K}$ , a negligible difference

As we discuss in the following section the difference in the characteristic dust temperatures of sources at low and high redshift may be due to a difference in the morphology of the dust emitting regions, as more extended sources will result in cooler temperatures, or, alternatively, due to the optical depth of the dust clouds. It is important to stress however that although the difference between the characteristic dust temperatures of the AS2UDS SMGs and low-redshift infrared-bright galaxies may appear modest, it is indicative of a significant difference between the properties of the dust emitting regions in these two galaxy populations, highlighting that SMGs should not be viewed as simple high redshift analogues of local U/LIRGs (see also Chapman et al. 2004; Pope et al. 2008; Hainline et al. 2009; Ivison et al. 2010; Rowlands et al. 2014).

### 5.3. Luminosity Density, Brightness Temperature and Optical Depth

A subset of 23 AS2UDS SMGs were detected at a sufficiently high SNR in our  $0.3''$ -resolution ALMA imaging to allow a measurement of the intrinsic sizes of their  $870\ \mu\text{m}$  emission regions (see Simpson et al. 2015a). The far-infrared luminosity surface density of the AS2UDS

SMGs as a function of the characteristic dust temperature of each SMG is shown in Figure 5, where we have assumed that half of the far-infrared luminosity is emitted within the half-light radius of the observed  $870\ \mu\text{m}$  emission. The data show a clear trend of increasing temperature with luminosity surface density, as expected if warmer dust emission traces regions of increasingly dense star formation.

However, we must consider that the observed peak of the far-infrared dust SED is sensitive to both the optical depth and the temperature of the dust. As such, the dust temperature derived from SED fitting is dependent on the frequency at which the optical depth is assumed to be unity. Instead, with measured sizes at  $870\ \mu\text{m}$ , we can determine the brightness temperature ( $T_B$ ), a fundamental property of the sources in our sample, by solving:

$$B_{\nu_{\text{rest}}}(T_B) = 0.5 S_{\nu_{\text{obs}}} (1+z)^3 / \Omega_{\nu_{\text{obs}}} \quad (4)$$

where  $\Omega_{\nu} = \pi R_{\nu}^2 / D_A^2$  and represents the solid angle subtended by the source,  $R_{\nu}$  is the intrinsic size of the emission region deconvolved from the beam, and a factor of 0.5 is included as we are considering the average brightness temperature within the half-light radius of each source. The brightness temperature represents the temperature of a blackbody radiating at a given luminosity and size. Using Equation 4 and the full expression for the Planck function we determine a median brightness temperature for the AS2UDS SMGs of  $25\ \text{K}$ , with a  $1\text{-}\sigma$  range of  $20\text{--}32\ \text{K}$  (Table 2).

The brightness temperature is a fundamental property of each source and can be related to the dust temperature and optical depth as follows

$$T_B = \frac{h\nu/k}{e^{h\nu/kT_D}}(1 - e^{-\tau_\nu}). \quad (5)$$

where  $\tau_\nu = (\frac{\nu}{\nu_0})^\beta$  and  $\beta = 1.8$  (see § 3.2.1). Thus, with resolved emission at multiple frequencies it is possible to determine the brightness temperature at different wavelengths and hence solve for both the true dust temperature and the optical depth (e.g. Sakamoto et al. 2008; Wilson et al. 2014; Barcos-Muñoz et al. 2015). We do not have resolved observations of the AS2UDS SMGs at multiple wavelengths, but we can use the measured size at an observed wavelength of  $870 \mu\text{m}$  and the far-infrared/sub-millimeter photometry from the *Herschel* and ALMA observations to place constraints on the optical depth. To do so, we assume that half-light radius of the observed emission at all far-infrared wavelengths is the same as that measured at  $870 \mu\text{m}$  (see also Aravena et al. 2008; Spilker et al. 2016). The emission from each SMG is modeled using Eqn. 5 and the fitting procedure described in § 3.2.1. We stress that for a given source the observed size of the emission region is dependent on the optical depth at the emission frequency; optical depth increases with frequency so shorter wavelength emission will appear extended on larger physical scales. As such, our assumption of a constant source size overestimates the fraction of observed 250, 350 and  $500 \mu\text{m}$  flux density that is located within the observed  $870\text{--}\mu\text{m}$  half-light radius, and the optical depth and the true dust temperature derived in our analysis should be considered lower limits.

To determine the optical depth of the SMGs we require that they are: resolved in our  $870 \mu\text{m}$  observations; detected in the SPIRE bands; and have a photometric redshift. 14 AS2UDS SMGs satisfy these criteria and we estimate that they have a median optical depth of unity at  $\lambda_0 (= c/\nu_0) \geq 75 \mu\text{m}$  ( $1\text{-}\sigma$  dispersion  $55\text{--}90 \mu\text{m}$ ),<sup>15</sup> and a true dust temperature of  $T_d > 33_{-2}^{+3} \text{K}$ , a systematic increase of  $\geq 3.1_{-0.3}^{+1.0} \text{K}$  relative to the median characteristic dust temperature that we that we measured for the same sources using an optically-thin SED model ( $T_d = 30 \pm 1 \text{K}$ , where the uncertainty represents the bootstrap error on the median of the sample).

With physically-motivated constraints on the dust temperature and the optical depth we can make an accurate comparison between the luminosity surface density and temperature relation that we determine for the AS2UDS SMGs and the Stefan-Boltzmann law. The Stefan-Boltzmann law between luminosity surface density and temperature is shown in Figure 5, and represents the expected relation for the faintest source in our sam-

ple ( $S_{870} = 3.6 \text{mJy}$ ) with a physical half-light radius of  $1.4 \text{kpc}$  (the median size of the sample; see Simpson et al. 2015a). While the AS2UDS SMGs follow the trend predicted by the Stefan-Boltzmann law they lie consistently below the predicted relation; an offset that increases linearly as we consider sources brighter than the faintest source in our sample. However, the Stefan-Boltzmann law is only valid for blackbody emission and can only be applied to SMGs if the dust emitting regions are assumed to be optically-thick at all wavelengths. Although, we have shown that the AS2UDS SMGs are optically-thick to at least  $75 \mu\text{m}$  we demonstrate that an increase in the optical depth raises the dust temperature of the sources an effect that only increases the offset between the Stefan-Boltzmann law and the SMGs presented here.

We now compare our optical depth measurements for the AS2UDS SMGs to similar studies of far-infrared-bright sources at high and low redshift. Recently, Riechers et al. (2013) used extensive FIR-photometry to place an upper limit of  $\lambda_0 < 163 \mu\text{m}$  and a best-fit value of  $\lambda_0 = 100_{-60}^{+40} \mu\text{m}$ , on the optical-depth of the high redshift SMG HFLS 3. While the result presented by Riechers et al. (2013) appears to be in agreement with the present work we caution that Riechers et al. (2013) consider the total emission from the source, whereas we have determined the typical optical depth within the half-light radius of the sub-millimeter emission. As seen for Arp220, a local ULIRG, the method presented here typically estimates a higher optical-depth for the dust emission when compared to fitting the global far-infrared photometry (see Rangwala et al. 2011; Scoville et al. 2017), consistent with the density of dust decreasing radially in these sources.

The 14 AS2UDS SMGs with optical depth constraints have a median  $L_{\text{FIR}} = 3.8_{-0.4}^{+1.1} \times 10^{12} L_\odot$  and so class as ULIRGs, despite our earlier analysis showing that they have distinct properties to these local sources. To investigate this discrepancy further we now compare our the optical depth of the AS2UDS SMGs to similarly luminous sources in the local Universe. Interferometric observations of CO or mm-emission have also been used to study selected  $z \approx 0$  ULIRGs (e.g. Scoville et al. 1991; Downes & Solomon 1998; Sakamoto et al. 1999; Bryant & Scoville 1999; Sakamoto et al. 2008). These studies have resulted in a consistent picture; that the emission from ULIRGs originates from a compact,  $R \lesssim 1 \text{kpc}$  region that has a high column density of molecular gas and is optically-thick at far-infrared wavelengths. For example, Scoville et al. (2017) presented high resolution ALMA observations of Arp 220 at  $2.6 \text{mm}$ . When combined with studies at shorter wavelengths the ALMA observations demonstrate that one of the two nuclei in Arp 220 is optically-thick to  $2.6 \text{mm}$ , a considerably higher optical depth than the lower limit that we have estimated for the AS2UDS SMGs.

Lutz et al. (2016) recently presented a study into the optical depth of 260 far-infrared luminous galaxies at  $z \lesssim 0.1$  with *Herschel*/PACS observations at  $70$  and  $160 \mu\text{m}$ . Restricting to sources with a comparable far-infrared luminosity to those studied here ( $L_{\text{FIR}} > 10^{12} L_\odot$ ), Lutz et al. (2016) show that far-infrared emission from these galaxies arises from regions that are, on average, optically-thin. This comparison

<sup>15</sup> To investigate whether the median optical depth of the AS2UDS SMGs is sensitive to our assumption of  $\beta = 1.8$  we repeat our analysis at a fixed value of  $\beta = 1.5$  and  $2.0$  (e.g. Magnelli et al. 2012; Dunne & Eales 2001). We find that the estimated optical depth of each SMG increases systematically with increasing  $\beta$ , and determine that the median wavelength at which the optical is unity is  $\geq 55 \mu\text{m}$  and  $\geq 85 \mu\text{m}$  for  $\beta = 1.5$  and  $2.0$ , respectively. Thus, we caution that our result is mildly sensitive to the assumed value of the dust emissivity and that we adopt a fixed value of  $\beta = 1.8$  throughout this work, consistent with previous studies.

suggests that the dust clouds in SMGs are optically-thick at longer wavelengths than local ULIRGs. However, Lutz et al. (2016) caution that the estimate of the optical-depth presented in their work should be considered a lower limit if the sources in their sample transition from optically-thick to -thin between 70 and 160  $\mu\text{m}$ . We have determined that the AS2UDS SMGs are optically-thick to  $\lambda_0 = 75_{-20}^{+15} \mu\text{m}$ , but similarly caution that this wavelength should be considered a lower limit. As such, it is clear that additional high resolution observations of the SMGs presented here, at multiple frequencies, are required to further investigate the difference in the optical depth properties of local ULIRGs and high redshift SMGs.

#### 5.4. Dust Correction

In §5.3 we showed that the dust emission region in SMGs becomes optically-thick, on average, at a wavelength of  $\lambda_0 \geq 75_{-20}^{+15} \mu\text{m}$ . The shape, and importantly, the peak wavelength of the far-infrared dust SED are sensitive to both the optical depth and temperature of the dust emission. We have modeled the far-infrared emission from each SMG with a physically-motivated model that includes an optical depth parameter and find that the dust temperatures are  $\geq 3.1_{-0.3}^{+1.0}$  K higher, on average, relative to the temperature derived from an optically-thin SED. A change in the dust temperature has a direct effect on the derived dust mass of each SMG, with an increase in the dust temperature resulting in a lower total dust mass. We derive dust masses for the 14 AS2UDS SMGs with optical depth constraints and find that they have a median dust mass of  $M_d = 6.3_{-0.8}^{+1.1} \times 10^8 M_\odot$ . In comparison, the median dust mass of these SMGs in the unphysical optically-thin regime is  $M_d = 8.5_{-0.5}^{+1.8} \times 10^8 M_\odot$ , which represents an increase of 35% relative to the optically-thick model. We note that when calculating dust mass we assume a dust opacity of  $\kappa_{850\mu\text{m}} = 0.07 \text{ m}^2 \text{ kg}^{-1}$  (James et al. 2002), but that there is likely a factor of three systematic uncertainty in this value (e.g. James et al. 2002; Alton et al. 2004). When combined with the half-light size of the dust emission the dust masses estimated from our optically-thick SED modeling imply that these SMGs have a median dust surface density of  $8.7_{-0.7}^{+1.3} \times 10^7 M_\odot \text{ kpc}^{-2}$ , within the half-light, or half-mass, radius of  $\sim 1\text{--}2 \text{ kpc}$ .

The high dust column densities that we have determined for the AS2UDS SMGs suggest that optical-to-near-infrared emission from the ongoing starburst will be strongly attenuated, comparable to that observed for local ULIRGs. As shown by Güver & Özel (2009), the magnitude of optical attenuation can be related to the column density of hydrogen atoms as follows;

$$N_H(\text{cm}^{-2}) = 2.21 \times 10^{21} A_V(\text{mag}) \quad (6)$$

where  $A_V$  represents extinction in the rest-frame  $V$ -band. To estimate the hydrogen column density of the AS2UDS SMGs we first convert the dust mass of each source to a gas mass by adopting a constant gas-to-dust ratio. We follow Swinbank et al. (2014), who present a comparison of the dust masses (Magnelli et al. 2012) and CO-derived gas masses (Bothwell et al. 2013) of the same SMGs and suggest that a gas-to-dust ratio

of  $\delta_{\text{gdr}} = 90 \pm 25$  is appropriate for SMGs; consistent with the expected gas-to-dust ratio given the metallicity, stellar mass and star formation rates of these sources (Draine et al. 2007; Rémy-Ruyer et al. 2014).

Adopting a gas-to-dust ratio of  $\delta_{\text{gdr}} = 90 \pm 25$  we find that the 14 SMGs in our resolved sample have a median hydrogen column density of  $N_H = 9.8_{-0.7}^{+1.4} \times 10^{23} \text{ cm}^{-2}$  and thus a median  $V$ -band dust obscuration of  $A_V = 540_{-40}^{+80} \text{ mag}$  to the source of the rest-frame  $\sim 200 \mu\text{m}$  emission. In our analysis we have assumed that the dust is uniformly distributed within the half-light radius, which is consistent with recent high resolution studies of non-lensed SMGs that indicate that the majority of the dust emission originates from a smooth “disk-like” component (see Hodge et al. 2016; Oteo et al. 2016; Gullberg et al. in prep.). We stress that the smooth appearance of the far-infrared emission does not rule out that the dust has an underlying “clumpy” morphology (e.g. Swinbank et al. 2010b; ALMA Partnership et al. 2015). Instead, we suggest that any “clumps” may be embedded in a dust photosphere that is optically-thick, with the far-infrared emission tracing the surface of the optically-thick region.

The average  $V$ -band extinction of  $540_{-40}^{+80} \text{ mag}$  for the AS2UDS SMGs indicates that effectively all of the optical-to-near-infrared emission from stars that are spatially-coincident with the far-infrared emission region will be attenuated by dust. Clearly the magnitude of the attenuation is extreme and we now consider the consequences for the multi-wavelength analysis of SMGs, and hence the physical properties that are determined from optical-to-near-infrared SED fitting (e.g. stellar mass and total star formation rate). In SED fitting routines the obscuration by dust is typically modelled by assuming that the dust is distributed in a uniform screen across the galaxy (e.g. Calzetti et al. 2000) or that the obscured star formation occurs in dense clouds embedded within the overall galaxy (e.g. MAGPHYS; da Cunha et al. 2008). A number of studies of interferometrically-identified SMGs have adopted these approaches, with both techniques finding that the detected visible emission from these typical SMGs has a median dust attenuation of  $A_V \sim 2 \text{ mag}$  (e.g. Simpson et al. 2014; da Cunha et al. 2015).

The magnitude of the  $V$ -band extinction in the AS2UDS SMGs is clearly in strong disagreement with estimates from optical-to-near-infrared SED fitting. To investigate this discrepancy we first consider the applicability of the simple dust screen model to the multi-wavelength analysis of SMGs. Using  $H$ -band *HST* imaging, Chen et al. (2015) showed that the stellar emission from SMGs has a median half-light radius of  $4.4_{-0.5}^{+1.1} \text{ kpc}$ , a factor of  $\sim 3\text{--}4 \times$  larger than the dust emission region (see Simpson et al. 2015a; Ikarashi et al. 2015) and hence an order of magnitude larger in area. Given the discrepancy in the profiles of the dust and the less-obscured stellar emission in SMGs it is unsurprising that a simple dust screen provides a poor representation of the resolved properties of these sources (see Hodge et al. 2016). Indeed, the discrepancy between the  $V$ -band attenuation estimated from the dust column density and the simple dust screen can be understood by considering that in SED-fitting the extinction is measured relative

to the detectable emission at restframe 1–2  $\mu\text{m}$  from a source. As such, the dust-correction is luminosity weighted by the light that is detectable from a source and in the scenario presented here should be considered as a lower limit for each SMG.

We next consider the MAGPHYS SED-fitting method presented by da Cunha et al. (2008). In MAGPHYS the dust emission from a source is considered to be comprised of two components; a diffuse interstellar medium and compact birth clouds. Energy-balance arguments are then employed to ensure that the emission in the far-infrared, which arises due to the reprocessing of stellar light by dust, is fully consistent with the integrated stellar light of the system. Thus, the energy-balance argument ensures that there is a physically-motivated upper limit to the total stellar emission of the system. However, we have shown that the far-infrared emission region in SMGs has a typical  $V$ -band obscuration of  $A_v = 540^{+80}_{-40}$  mag, which corresponds to an optical depth of  $\tau_v \sim 500^{+60}_{-30}$ . This optical depth is an order of magnitude higher than the model values used in MAGPHYS, and three orders of magnitude higher than the relatively tight prior that is placed on the parameter values (e.g. da Cunha et al. 2015).

Indeed, the results presented here suggest that any “energy-balance” analysis of the emission from SMGs should take into account the spatial distribution of the *detectable* optical and far-infrared components; the visible component of the emission and the far-infrared emission by dust should be considered as arising from effectively independent regions of the overall system. However, we strongly caution that in any “energy-balance” analysis of SMGs the detectable emission at optical wavelengths will contain no information about the ongoing obscured starburst. Thus, while an “energy-balance” analysis will determine a stellar component (including stellar mass) that is consistent with the far-infrared emission we stress that it is determined purely by the prior assumptions on the model parameters. As such, the stellar masses of these systems should be treated with extreme caution (see also Hainline et al. 2011; Michałowski et al. 2012; Simpson et al. 2014).

### 5.5. Spheroid Growth?

The intensity of the ongoing starburst in SMGs, along with their large gas reservoirs that can sustain a prolonged period of stellar mass growth ( $\sim 100$  Myr; Bothwell et al. 2013), has led a number of authors to suggest an evolutionary link between these sources and local elliptical galaxies (e.g. Lilly et al. 1999; Blain et al. 2004; Swinbank et al. 2006; Tacconi et al. 2008; Swinbank et al. 2010a; Simpson et al. 2014; Toft et al. 2014). Furthermore, the recent discovery that the ongoing obscured star formation in SMGs has a half-light radius of  $\sim 1.5$  kpc has fueled speculation that we are witnessing the direct assembly of a spheroid component in these sources (Simpson et al. 2015a; Ikarashi et al. 2015; Hodge et al. 2016).

To investigate the possible descendants of the AS2UDS we now estimate the final stellar mass that will be contained within the compact, starburst region. The stellar mass of the ongoing starbursts cannot be estimated through SED-fitting techniques due to the ab-

sence of detectable emission at optical-to-near-infrared wavelengths. However, we can constrain the final stellar mass component of the ongoing starburst by considering the gas masses derived for the AS2UDS SMGs with measured sizes at restframe  $\sim 200$   $\mu\text{m}$ . Assuming that all of the available gas in these SMGs is converted into stars then the ongoing starburst will result in a compact spheroid with a mass of  $\sim 1 \times 10^{11} M_\odot$ . The assumption that all of the gas mass is converted to stellar mass is unrealistic, with large-scale outflows likely removing some gas from the galaxy. However, we have not included any contribution to the final stellar mass from either a pre-existing stellar component or the transformation of the extended stellar component due to a potential ongoing merger in these systems (Chen et al. 2015). As such, we consider  $\sim 10^{11} M_\odot$  as a reasonable estimate of the stellar mass of the post-starburst remnant.

In the local Universe, the high-mass end of the galaxy stellar mass function ( $\gtrsim 1 \times 10^{11} M_\odot$ ) is dominated by elliptical, lenticular (S0), and, to a significantly lesser extent, massive (Sa) spiral galaxies (see Kelvin et al. 2014). To determine a feasible evolutionary pathway for the AS2UDS SMGs we first investigate whether the estimated spheroid masses are consistent with the properties of local S0 and Sa galaxies, which have typical a spheroid-to-disc mass ratio of  $\sim 0.4$  (Graham & Worley 2008). We have estimated that the AS2UDS SMGs will form a spheroid component with a stellar mass of  $\sim 10^{11} M_\odot$ . Thus, if the AS2UDS SMGs are the progenitors of S0 and Sa galaxies then they must correspond to galaxies in the local Universe that have total stellar masses of  $\sim 2\text{--}3 \times 10^{11} M_\odot$ , where we have neglected to include any stellar mass growth between the SMG-phase and  $z=0$ .

S0 and Sa galaxies with total stellar masses of  $\sim 2\text{--}3 \times 10^{11} M_\odot$  are extremely rare, with an estimated space density of  $\sim 10^{-7}\text{--}10^{-8} \text{Mpc}^{-3}$ <sup>16</sup> (Kelvin et al. 2014). The AS2UDS SMG considered here have a median 870  $\mu\text{m}$  flux density of  $8.0 \pm 0.4$  mJy, corresponding to an estimated comoving space density of  $\sim 10^{-5} \text{Mpc}^{-3}$  (Karim et al. 2013; Simpson et al. 2014, 2015b). Thus, these SMGs are expected to be 2–3 orders of magnitude more numerous than local spirals (Sa) and lenticular galaxies that have spheroid stellar masses that are consistent with these high redshift starbursts. As such, we suggest that *if* we are indeed witnessing a centrally-concentrated starburst in SMGs, which is directly growing the stellar mass of a spheroidal component, then they cannot evolve into local spiral or lenticular galaxies (without subsequent significant loss of stellar mass). Instead, SMGs must be the progenitors of today’s massive ellipticals and hence are ideal tracers for the formation of the most massive and oldest galaxies at high redshift (Nelán et al. 2005).

## 6. CONCLUSIONS

In this paper we have presented a comprehensive study of the multi-wavelength properties of 52 ALMA-identified SMGs in the UDS field with high spatial res-

<sup>16</sup> We estimate the comoving space density of local S0 & Sa galaxies by integrating the best-fit Schechter functions to the morphological-type stellar mass functions presented by Kelvin et al. (2014), which were derived from observations taken as part of the Galaxy and Mass Assembly survey (GAMA)

olution 870  $\mu\text{m}$  imaging. The main conclusions of our work are:

- We use the available imaging of the UDS to characterize the properties of the AS2UDS SMGs and show that 35 of the 48 (non-lensed) AS2UDS SMGs ( $73 \pm 7\%$ ) are detected in the  $K$ -band imaging. We estimate photometric redshifts for these 35 AS2UDS SMGs and determine that they lie at a median redshift of  $z = 2.65 \pm 0.13$ , which rises to  $z = 2.9 \pm 0.2$  if the SMGs with insufficient photometry to derive a photometric redshift are included and are assumed to lie uniformly at  $z = 3-6$ .
- We model the far-infrared emission from the sources in our sample and show that they have a median far-infrared luminosity of  $L_{\text{FIR}} = 3.2^{+0.3}_{-0.7} \times 10^{12} L_{\odot}$ . Combining our AS2UDS sample with the ALESS survey we find that SMGs are  $\sim 8$  K cooler, at a fixed far-infrared luminosity, compared to local far-infrared-bright galaxies. This is consistent with the larger physical size of the high redshift sources and suggests that SMGs should not be considered scaled-up version of  $z \sim 0$  ULIRGs.
- We use a subset of 23 AS2UDS SMGs that have dust emission sizes from high resolution ALMA imaging to constraint the optical depth of the SMG population. The far-infrared sizes show that this emission does not originate from dust clouds that are optically-thin. Instead, we show that the the dust regions in these archetypal SMGs are optically-thick at a median wavelength of  $\lambda_0 \geq 75 \mu\text{m}$ , with a  $1-\sigma$  dispersion of  $55-90 \mu\text{m}$ . By modeling the emission with an optically-thick SED we estimate these sources have a true dust temperature that is  $\geq 3.1^{+1.0}_{-0.3}$  K higher than the characteristic dust temperature that is measured by assuming the emission is optically-thin at all wavelengths. Thus, the discrepancy in the characteristic dust temperatures of SMGs and local ULIRGs may be due to SMGs being more optically-thick and larger, or intrinsically cooler and larger.
- Using the dust masses derived from our physically-motivated, optically-thick SED fits we determine that AS2UDS SMGs have a median hydrogen column density of  $N_H = 9.8^{+1.4}_{-0.7} \times 10^{23} \text{cm}^{-2}$ , corresponding to a median  $V$ -band obscuration of  $A_v = 540^{+80}_{-40}$  mag averaged along the line of sight to the source of the far-infrared emission. The extreme attenuation in the far-infrared emission region means that effectively all of the stellar light

from any co-located stellar component is obscured at optical-to-near-infrared wavelengths. As such, stellar properties that are derived through SED-fitting techniques should be treated with the utmost caution.

- Finally, we investigate the possible evolutionary pathways for the AS2UDS SMGs. Assuming that the compact, obscured starburst is centrally-concentrated, we estimate that the AS2UDS SMGs will host a post-starburst spheroid with a stellar mass of  $\sim 10^{11} M_{\odot}$ . We show that local S0 & Sa galaxies with a comparable spheroidal mass have a space density that is 2-3 orders of magnitude lower than the AS2UDS SMGs, indicating that SMGs do not evolve into lenticular or spiral galaxies. Instead, our analysis indicates that SMGs must be the progenitors of local elliptical galaxies.

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All data used in this analysis can be obtained from either the ALMA archive, the Canadian Data Archive Center (CADCA/JCMT), or the WFCAM Science Archive (WSA/UKIDSS).

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TABLE 1: OBSERVED PROPERTIES

ID	R.A. (J2000)	Dec. (J2000)	$K^b$ (AB)	$S_{250}$ (mJy)	$S_{350}$ (mJy)	$S_{500}$ (mJy)	$S_{870}^{\text{ALMA}}$ (mJy)	$S_{1.4\text{ GHz}}$ ( $\mu\text{Jy}$ )
UDS47.0	02:19:24.84	-05:09:20.7	< 24.6	< 9.2	< 10.6	< 12.2	$8.7 \pm 0.6$	$85 \pm 21$
UDS47.1	02:19:24.64	-05:09:16.3	< 24.6	< 9.2	< 10.6	< 12.2	$2.1 \pm 0.8$	-
UDS48.0	02:19:24.57	-04:53:00.2	$21.49 \pm 0.02$	$85.2 \pm 7.8$	$64.5 \pm 6.7$	$26.3 \pm 5.1$	$7.5 \pm 0.5$	$254 \pm 22$
UDS48.1	02:19:24.62	-04:52:56.9	$22.37 \pm 0.05$	< 18.1	< 17.0	< 17.8	$1.3 \pm 0.5$	$67 \pm 20$
UDS57.0	02:19:21.14	-04:56:51.3	$22.40 \pm 0.05$	< 16.7	< 16.5	< 18.6	$9.5 \pm 0.6$	$65 \pm 21$
UDS57.1	02:19:20.88	-04:56:52.9	< 24.6	$27.9 \pm 4.2$	$36.3 \pm 5.3$	$37.2 \pm 6.4$	$5.9 \pm 0.9$	-
UDS57.2	02:19:21.41	-04:56:49.0	$25.08 \pm 0.45$	< 14.1	< 14.9	< 18.6	$1.5 \pm 0.6$	-
UDS57.3	02:19:21.39	-04:56:38.8	< 24.6	< 12.3	< 13.9	< 17.4	$2.1 \pm 1.0$	-
UDS74.0	02:19:13.19	-04:47:08.0	$22.53 \pm 0.05$	< 7.7	$20.1 \pm 3.9$	$19.4 \pm 4.1$	$4.5 \pm 0.5$	-
UDS74.1	02:19:13.19	-04:47:05.0	$24.24 \pm 0.23$	< 9.0	< 10.8	< 13.9	$1.5 \pm 0.5$	-
UDS78.0	02:19:09.74	-05:15:30.6	$22.82 \pm 0.08$	$27.3 \pm 4.1$	$30.7 \pm 4.9$	$21.6 \pm 4.4$	$8.2 \pm 0.5$	$63 \pm 22$
UDS79.0	02:19:09.94	-05:00:08.6	$22.99 \pm 0.07$	< 8.5	$16.2 \pm 3.5$	$14.8 \pm 3.4$	$7.7 \pm 0.5$	$65 \pm 17$
UDS109.0 <sup>a</sup>	02:18:50.07	-05:27:25.5	-	< 9.2	< 15.5	< 13.9	$7.6 \pm 0.7$	$131.5 \pm 31.8$
UDS109.1	02:18:50.30	-05:27:17.2	$22.23 \pm 0.04$	$11.4 \pm 2.4$	$24.2 \pm 4.4$	$25.5 \pm 5.0$	$4.2 \pm 0.6$	-
UDS110.0	02:18:48.24	-05:18:05.2	$21.17 \pm 0.02$	$27.0 \pm 4.1$	$26.4 \pm 4.6$	$18.6 \pm 4.0$	$7.7 \pm 0.6$	$125 \pm 18$
UDS110.1	02:18:48.76	-05:18:02.1	$21.20 \pm 0.02$	$20.2 \pm 3.5$	$20.4 \pm 4.0$	< 16.0	$2.0 \pm 0.8$	-
UDS156.0	02:18:24.14	-05:22:55.3	$23.09 \pm 0.09$	< 17.8	< 17.0	< 18.6	$9.7 \pm 0.7$	$39.0 \pm 11.2$
UDS156.1	02:18:24.24	-05:22:56.9	$24.10 \pm 0.21$	$33.0 \pm 4.6$	$34.6 \pm 5.2$	$36.5 \pm 6.3$	$8.5 \pm 0.7$	$136 \pm 45$
UDS160.0 <sup>a</sup>	02:18:23.73	-05:11:38.5	-	$16.5 \pm 3.1$	$20.6 \pm 4.0$	$13.0 \pm 3.1$	$7.9 \pm 0.6$	$44 \pm 8$
UDS168.0	02:18:20.40	-05:31:43.2	$21.96 \pm 0.04$	< 12.3	$17.7 \pm 3.7$	$16.1 \pm 3.7$	$6.7 \pm 0.6$	$135 \pm 32$
UDS168.1	02:18:20.31	-05:31:41.7	$21.96 \pm 0.04$	$18.3 \pm 3.2$	< 16.3	< 16.6	$2.7 \pm 0.6$	-
UDS168.2	02:18:20.17	-05:31:38.6	< 24.6	< 11.1	< 16.3	< 16.6	$1.5 \pm 0.7$	-
UDS199.0	02:18:07.18	-04:44:13.8	< 24.6	< 9.2	< 10.8	< 12.2	$4.2 \pm 0.6$	$68 \pm 19$
UDS199.1	02:18:07.19	-04:44:10.9	$25.36 \pm 0.60$	< 9.2	< 10.8	< 12.2	$2.4 \pm 0.5$	-
UDS202.0	02:18:05.65	-05:10:49.6	$23.89 \pm 0.16$	$13.0 \pm 2.6$	$22.8 \pm 4.2$	$18.3 \pm 4.0$	$10.5 \pm 0.5$	$72 \pm 16$
UDS202.1	02:18:05.05	-05:10:46.3	$24.27 \pm 0.22$	< 8.1	< 9.9	< 13.9	$3.5 \pm 0.9$	$48 \pm 7$
UDS204.0	02:18:03.01	-05:28:41.9	$23.74 \pm 0.16$	< 8.1	$12.6 \pm 3.0$	< 12.9	$11.6 \pm 0.6$	$74 \pm 22$
UDS204.1	02:18:03.01	-05:28:32.5	< 24.6	< 9.2	< 10.8	< 14.5	$2.2 \pm 0.9$	-
UDS216.0	02:17:56.74	-04:52:38.9	$21.01 \pm 0.01$	$23.6 \pm 3.8$	$24.6 \pm 4.4$	$14.5 \pm 3.4$	$5.2 \pm 0.5$	$88 \pm 17$
UDS218.0	02:17:54.80	-05:23:23.0	$22.63 \pm 0.06$	$16.0 \pm 2.9$	$15.3 \pm 3.4$	$21.1 \pm 4.4$	$6.6 \pm 0.7$	$58 \pm 18$
UDS269.0 <sup>a</sup>	02:17:30.44	-05:19:22.4	-	< 10.0	$12.8 \pm 3.1$	$23.2 \pm 4.7$	$12.9 \pm 0.6$	$46 \pm 15$
UDS269.1	02:17:30.25	-05:19:18.4	$22.33 \pm 0.05$	$12.1 \pm 2.5$	< 14.6	< 16.8	$2.0 \pm 0.7$	-
UDS286.0 <sup>a</sup>	02:17:25.73	-05:25:41.2	-	$12.1 \pm 2.5$	$15.0 \pm 3.4$	< 18.3	$5.1 \pm 0.7$	$103 \pm 19$
UDS286.1	02:17:25.63	-05:25:33.7	$23.95 \pm 0.20$	< 17.6	< 16.2	< 18.3	$5.0 \pm 0.6$	-
UDS286.2	02:17:25.80	-05:25:37.5	< 24.6	$14.1 \pm 2.8$	$17.5 \pm 3.7$	$16.0 \pm 3.7$	$2.6 \pm 0.6$	-
UDS286.3	02:17:25.52	-05:25:36.7	< 24.6	< 17.6	< 15.7	< 18.3	$1.4 \pm 0.6$	-
UDS292.0	02:17:21.53	-05:19:07.8	$22.35 \pm 0.04$	$17.2 \pm 3.2$	$13.1 \pm 3.1$	$17.4 \pm 3.9$	$4.0 \pm 0.8$	$52 \pm 17$
UDS292.1	02:17:21.96	-05:19:09.8	$21.93 \pm 0.03$	$17.9 \pm 3.2$	$19.8 \pm 3.9$	< 15.7	$3.6 \pm 0.8$	-
UDS298.0	02:17:19.57	-05:09:41.2	$21.83 \pm 0.03$	$13.9 \pm 2.7$	$12.7 \pm 3.0$	< 13.9	$1.3 \pm 0.4$	-
UDS298.1	02:17:19.46	-05:09:33.2	$22.05 \pm 0.03$	< 10.0	< 12.6	< 12.2	$1.6 \pm 0.8$	-
UDS306.0	02:17:17.07	-05:33:26.6	$21.22 \pm 1.73$	$53.2 \pm 6.2$	$36.7 \pm 5.4$	< 16.3	$8.3 \pm 0.5$	$95 \pm 22$
UDS306.1	02:17:17.16	-05:33:32.5	$21.31 \pm 1.89$	$42.4 \pm 5.5$	$30.4 \pm 4.9$	$29.1 \pm 5.4$	$2.4 \pm 0.4$	$224 \pm 30$
UDS306.2	02:17:16.81	-05:33:31.8	< 24.6	< 18.1	< 16.5	< 17.2	$2.3 \pm 0.9$	-
UDS334.0	02:17:02.47	-04:57:20.0	$21.49 \pm 0.02$	$34.6 \pm 4.8$	$26.7 \pm 4.6$	$15.9 \pm 3.6$	$3.6 \pm 0.8$	$783 \pm 16$
UDS345.0	02:16:57.61	-05:20:38.6	$21.47 \pm 0.02$	$18.0 \pm 3.2$	$24.5 \pm 4.4$	< 15.5	$2.0 \pm 0.7$	$74 \pm 21$
UDS361.0	02:16:47.92	-05:01:29.8	$22.02 \pm 0.03$	$14.1 \pm 2.8$	$27.8 \pm 4.7$	$23.1 \pm 4.6$	$11.8 \pm 0.6$	$68 \pm 22$
UDS361.1	02:16:47.73	-05:01:25.8	$23.64 \pm 0.15$	< 9.0	< 13.9	< 14.8	$2.0 \pm 0.7$	-
UDS377.0	02:16:41.11	-05:03:51.4	< 24.6	$14.7 \pm 2.9$	$16.2 \pm 3.5$	< 15.7	$8.1 \pm 0.5$	-
UDS392.0	02:16:33.29	-05:11:59.0	$23.71 \pm 0.14$	< 9.2	< 11.2	< 12.2	$3.7 \pm 0.5$	-
UDS408.0	02:16:22.26	-05:11:07.8	$22.15 \pm 0.04$	$20.8 \pm 3.6$	< 15.9	< 13.9	$9.1 \pm 0.7$	$101 \pm 20$
UDS408.1	02:16:22.28	-05:11:11.9	< 24.6	< 10.6	$20.6 \pm 4.0$	$15.8 \pm 3.6$	$2.1 \pm 0.9$	-
UDS412.0	02:16:20.13	-05:17:26.2	< 24.6	$15.4 \pm 2.9$	$26.3 \pm 4.5$	$19.5 \pm 4.1$	$6.6 \pm 0.7$	-

<sup>a</sup> Identified as a potentially lensed SMG, <sup>b</sup> Total magnitude.

TABLE 2: PHYSICAL PROPERTIES

ID	$z_{phot}$	$L_{FIR}^{Thin, b}$ ( $\times 10^{12} L_{\odot}$ )	$T_d^{Thin, b}$ (K)	FWHM <sup>c</sup> (")	$T_B^d$ (K)	$L_{FIR}^{Thick, e}$ ( $\times 10^{12} L_{\odot}$ )	$T_d^{Thick, e}$ (K)	$\lambda_0^e$ $\mu m$
UDS47.0	-	-	-	0.28 ± 0.03	-	-	-	-
UDS47.1	-	-	-	-	-	-	-	-
UDS48.0	2.14 <sup>+0.07</sup> <sub>-0.15</sub>	11.81 <sup>+1.25</sup> <sub>-2.26</sub>	39.7 <sup>+1.3</sup> <sub>-2.4</sub>	0.28 ± 0.02	24.4	11.14 <sup>+1.40</sup> <sub>-1.82</sub>	45.7 <sup>+1.7</sup> <sub>-2.3</sub>	74 <sup>+7</sup> <sub>-7</sub>
UDS48.1	2.25 <sup>+0.15</sup> <sub>-0.15</sub>	0.05–3.08	< 44	-	-	-	-	-
UDS57.0	1.87 <sup>+0.24</sup> <sub>-0.17</sub>	0.23–0.93	< 18	0.34 ± 0.02	20.3	-	-	-
UDS57.1	-	-	-	0.26 ± 0.05	-	-	-	-
UDS57.2	2.65 <sup>+0.22</sup> <sub>-0.31</sub>	0.09–3.06	< 43	-	-	-	-	-
UDS57.3	-	-	-	-	-	-	-	-
UDS74.0	3.26 <sup>+0.05</sup> <sub>-0.11</sub>	3.61 <sup>+0.43</sup> <sub>-0.82</sub>	33.5 <sup>+1.3</sup> <sub>-2.6</sub>	0.38 ± 0.04	24.6	3.60 <sup>+0.51</sup> <sub>-0.70</sub>	36.2 <sup>+1.7</sup> <sub>-2.2</sub>	56 <sup>+9</sup> <sub>-7</sub>
UDS74.1	4.32 <sup>+0.37</sup> <sub>-0.83</sub>	0.27–5.85	< 55	-	-	-	-	-
UDS78.0	2.80 <sup>+0.22</sup> <sub>-0.19</sub>	5.98 <sup>+0.90</sup> <sub>-1.11</sub>	33.1 <sup>+1.5</sup> <sub>-2.0</sub>	0.35 ± 0.03	27.9	5.86 <sup>+1.07</sup> <sub>-0.91</sub>	37.4 <sup>+2.1</sup> <sub>-1.9</sub>	77 <sup>+8</sup> <sub>-7</sub>
UDS79.0	3.27 <sup>+0.07</sup> <sub>-0.31</sub>	3.43 <sup>+0.21</sup> <sub>-1.03</sub>	28.6 <sup>+0.6</sup> <sub>-3.2</sub>	0.43 ± 0.02	27.4	3.36 <sup>+0.39</sup> <sub>-0.83</sub>	31.6 <sup>+1.1</sup> <sub>-2.7</sub>	78 <sup>+7</sup> <sub>-11</sub>
UDS109.0 <sup>a</sup>	-	-	-	-	-	-	-	-
UDS109.1	2.65 <sup>+0.19</sup> <sub>-0.09</sub>	2.88 <sup>+0.50</sup> <sub>-0.40</sub>	31.3 <sup>+1.8</sup> <sub>-1.5</sub>	-	-	-	-	-
UDS110.0	1.68 <sup>+0.24</sup> <sub>-0.10</sub>	1.60 <sup>+0.35</sup> <sub>-0.31</sub>	23.5 <sup>+1.4</sup> <sub>-1.4</sub>	0.28 ± 0.02	19.7	1.59 <sup>+0.42</sup> <sub>-0.26</sub>	27.5 <sup>+2.0</sup> <sub>-1.3</sub>	125 <sup>+16</sup> <sub>-13</sub>
UDS110.1	2.80 <sup>+0.04</sup> <sub>-0.07</sub>	5.26 <sup>+1.39</sup> <sub>-1.12</sub>	44.7 <sup>+5.5</sup> <sub>-4.6</sub>	-	-	-	-	-
UDS156.0	3.67 <sup>+0.12</sup> <sub>-0.13</sub>	1.23–4.74	< 29	0.25 ± 0.02	55.5	-	-	-
UDS156.1	2.35 <sup>+0.57</sup> <sub>-0.26</sub>	4.83 <sup>+1.50</sup> <sub>-1.47</sub>	30.3 <sup>+2.6</sup> <sub>-3.0</sub>	0.24 ± 0.03	32.4	4.84 <sup>+2.04</sup> <sub>-1.17</sub>	37.9 <sup>+4.5</sup> <sub>-3.1</sub>	120 <sup>+20</sup> <sub>-16</sub>
UDS160.0 <sup>a</sup>	-	-	-	-	-	-	-	-
UDS168.0	2.77 <sup>+0.06</sup> <sub>-0.17</sub>	2.65 <sup>+0.33</sup> <sub>-0.71</sub>	27.5 <sup>+1.1</sup> <sub>-2.7</sub>	0.42 ± 0.03	22.3	2.63 <sup>+0.41</sup> <sub>-0.60</sub>	30.1 <sup>+1.5</sup> <sub>-2.2</sub>	74 <sup>+9</sup> <sub>-10</sub>
UDS168.1	2.77 <sup>+0.06</sup> <sub>-0.17</sub>	3.95 <sup>+0.70</sup> <sub>-1.02</sub>	40.3 <sup>+3.3</sup> <sub>-4.1</sub>	-	-	-	-	-
UDS168.2	-	-	-	-	-	-	-	-
UDS199.0	-	-	-	0.28 ± 0.06	-	-	-	-
UDS199.1	5.01 <sup>+0.37</sup> <sub>-2.01</sub>	0.62–8.34	< 55	-	-	-	-	-
UDS202.0	3.62 <sup>+0.44</sup> <sub>-0.28</sub>	7.06 <sup>+1.16</sup> <sub>-1.44</sub>	32.8 <sup>+1.7</sup> <sub>-2.4</sub>	0.36 ± 0.02	39.9	7.03 <sup>+1.51</sup> <sub>-1.19</sub>	38.6 <sup>+2.6</sup> <sub>-2.2</sub>	89 <sup>+8</sup> <sub>-9</sub>
UDS202.1	3.35 <sup>+0.66</sup> <sub>-0.35</sub>	0.36–2.55	< 33	-	-	-	-	-
UDS204.0	3.44 <sup>+0.59</sup> <sub>-0.21</sub>	3.33 <sup>+0.78</sup> <sub>-0.93</sub>	24.5 <sup>+1.8</sup> <sub>-3.0</sub>	0.58 ± 0.02	26.9	3.32 <sup>+1.12</sup> <sub>-0.60</sub>	27.2 <sup>+2.8</sup> <sub>-1.7</sub>	89 <sup>+12</sup> <sub>-12</sub>
UDS204.1	-	-	-	-	-	-	-	-
UDS216.0	2.19 <sup>+0.05</sup> <sub>-0.09</sub>	2.84 <sup>+0.33</sup> <sub>-0.50</sub>	30.3 <sup>+1.3</sup> <sub>-1.9</sub>	0.70 ± 0.04	12.6	2.80 <sup>+0.41</sup> <sub>-0.42</sub>	31.0 <sup>+1.7</sup> <sub>-1.5</sub>	32 <sup>+4</sup> <sub>-3</sub>
UDS218.0	3.00 <sup>+0.17</sup> <sub>-0.25</sub>	4.02 <sup>+0.48</sup> <sub>-0.89</sub>	31.5 <sup>+1.3</sup> <sub>-2.5</sub>	0.37 ± 0.04	26.3	3.94 <sup>+0.62</sup> <sub>-0.75</sub>	34.7 <sup>+2.0</sup> <sub>-2.1</sub>	70 <sup>+11</sup> <sub>-9</sub>
UDS269.0 <sup>a</sup>	-	-	-	-	-	-	-	-
UDS269.1	2.61 <sup>+0.26</sup> <sub>-0.10</sub>	2.37 <sup>+0.76</sup> <sub>-0.61</sub>	37.9 <sup>+6.0</sup> <sub>-4.2</sub>	-	-	-	-	-
UDS286.0 <sup>a</sup>	-	-	-	-	-	-	-	-
UDS286.1	4.91 <sup>+0.20</sup> <sub>-0.76</sub>	1.26–10.55	< 47	0.26 ± 0.07	54.8	-	-	-
UDS286.2	-	-	-	-	-	-	-	-
UDS286.3	-	-	-	-	-	-	-	-
UDS292.0	2.65 <sup>+0.25</sup> <sub>-0.07</sub>	3.09 <sup>+0.59</sup> <sub>-0.57</sub>	33.5 <sup>+2.6</sup> <sub>-2.7</sub>	-	-	-	-	-
UDS292.1	2.51 <sup>+0.23</sup> <sub>-0.10</sub>	3.03 <sup>+0.80</sup> <sub>-0.48</sub>	34.2 <sup>+3.5</sup> <sub>-2.4</sub>	-	-	-	-	-
UDS298.0	1.81 <sup>+0.20</sup> <sub>-0.10</sub>	1.24 <sup>+0.52</sup> <sub>-0.32</sub>	34.1 <sup>+4.9</sup> <sub>-3.5</sub>	-	-	-	-	-
UDS298.1	2.01 <sup>+0.21</sup> <sub>-0.18</sub>	0.05–1.04	< 31	-	-	-	-	-
UDS306.0	2.31 <sup>+0.06</sup> <sub>-0.21</sub>	6.39 <sup>+0.53</sup> <sub>-1.39</sub>	33.6 <sup>+0.9</sup> <sub>-2.3</sub>	0.30 ± 0.02	26.1	6.15 <sup>+0.70</sup> <sub>-1.13</sub>	38.7 <sup>+1.3</sup> <sub>-2.2</sub>	86 <sup>+7</sup> <sub>-7</sub>
UDS306.1	1.28 <sup>+0.53</sup> <sub>-0.06</sub>	2.15 <sup>+0.84</sup> <sub>-0.49</sub>	32.9 <sup>+3.8</sup> <sub>-2.5</sub>	-	-	-	-	-
UDS306.2	-	-	-	-	-	-	-	-
UDS334.0	1.93 <sup>+0.08</sup> <sub>-0.17</sub>	3.43 <sup>+0.53</sup> <sub>-0.94</sub>	34.7 <sup>+2.3</sup> <sub>-3.6</sub>	-	-	-	-	-
UDS345.0	1.69 <sup>+0.26</sup> <sub>-0.05</sub>	1.36 <sup>+0.47</sup> <sub>-0.24</sub>	30.0 <sup>+3.8</sup> <sub>-2.4</sub>	-	-	-	-	-
UDS361.0	3.08 <sup>+0.18</sup> <sub>-0.29</sub>	5.51 <sup>+0.60</sup> <sub>-1.16</sub>	29.0 <sup>+1.0</sup> <sub>-2.1</sub>	0.62 ± 0.02	23.2	5.47 <sup>+0.77</sup> <sub>-0.98</sub>	31.3 <sup>+1.4</sup> <sub>-1.8</sub>	63 <sup>+5</sup> <sub>-5</sub>
UDS361.1	0.61 <sup>+0.04</sup> <sub>-0.11</sub>	0.00–0.05	< 15	-	-	-	-	-
UDS377.0	-	-	-	0.16 ± 0.02	-	-	-	-
UDS392.0	1.72 <sup>+1.57</sup> <sub>-0.06</sub>	0.07–0.58	< 21	< 0.18	> 22	-	-	-
UDS408.0	2.62 <sup>+0.05</sup> <sub>-0.13</sub>	3.30 <sup>+0.34</sup> <sub>-0.78</sub>	27.4 <sup>+1.0</sup> <sub>-2.4</sub>	0.66 ± 0.04	17.6	3.29 <sup>+0.43</sup> <sub>-0.66</sub>	28.8 <sup>+1.3</sup> <sub>-1.9</sub>	52 <sup>+6</sup> <sub>-7</sub>
UDS408.1	-	-	-	-	-	-	-	-
UDS412.0	-	-	-	0.30 ± 0.07	-	-	-	-

<sup>a</sup> Identified as a potentially lensed SMG, <sup>b</sup> Assuming an optically–thin SED. The full range of plausible values are given for sources that are only detected in the far-infrared at 870  $\mu m$ . <sup>c</sup> Intrinsic source size, corrected for synthesized beam, at observed 870  $\mu m$  (see Simpson et al. 2015a) <sup>d</sup> Average brightness temperature of the dust contained within the half–light radius of the observed 870  $\mu m$  emission <sup>e</sup> Assuming an optically–thick SED, and using observed size of the 870  $\mu m$  emission as a Gaussian prior in the FIR SED fitting.