The X-ray Properties of Weak Lensing Selected Galaxy Clusters

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ABSTRACT

We present the results of an X-ray follow-up campaign targeting 10 Weak Lensing (WL) selected galaxy clusters from a Subaru WL survey. Eight clusters were studied with dedicated Chandra pointings, whereas archival X-ray data were used for the remaining two clusters. The WL clusters appear to fit the same scaling relation between X-ray luminosity and temperature as X-ray selected clusters. However, when we consider the luminosity-mass relation, the WL selected clusters appear underluminous by a factor $3.9 \pm 0.9$ (or, alternatively, more massive by $2.9 \pm 0.2$), compared to X-ray selected clusters. Only by considering various observational effects that could potentially bias WL masses, can this difference be reconciled. We used X-ray imaging data to quantify the dynamical state of the clusters and found that one of the clusters appears dynamically relaxed, and two of the clusters host a cool core, consistent with Sunyaev-Zel’dovich effect selected clusters. Our results suggest that regular, cool core clusters may be over-represented in X-ray selected samples.

Key words: Galaxies: clusters: general – Cosmology: observations – X-rays: galaxies: clusters – Gravitational lensing

1 INTRODUCTION

Clusters of galaxies are the largest gravitationally-collapsed structures in the Universe. The study of their number density and growth from high density perturbations in the early Universe offers insight into the underlying cosmology (e.g., Mantz et al. 2008; Vikhlinin et al. 2009b). However, the use of clusters as a cosmological probe requires an efficient method to find clusters over a wide redshift range, an observational method of determining the cluster mass and a method to compute the selection function or the survey volume in which clusters are found (Mantz et al. 2008). A variety of observational techniques has been brought to bear on these requirements, each with different strengths and weaknesses.

Early samples of clusters were based upon optical selection, however these samples have traditionally suffered from projection effects and uncertainties in the optical richness-mass relation. Recently, the first clusters have been detected in blind surveys using the Sunyaev-Zel’dovich effect (SZE; Staniszewski et al. 2009; Vanderlinde et al. 2010; Planck Collaboration et al. 2011a). This method holds much promise due to the redshift independence of the SZE (Birkinshaw 1999), however the use of the SZE as a mass estimator remains largely untested. Until now, the most effective method of building large, well defined cluster samples has been via X-ray selection. The high X-ray luminosities of clusters makes it relatively easy to detect and study clusters out to high redshifts, and X-ray cluster studies have provided a means to impose tight constraints on cosmological parameters (e.g., Vikhlinin et al. 2009b; Mantz et al. 2010b).

A common weakness of the three techniques outlined above is that the clusters are selected based on the properties of their minority baryonic content. However, cosmological constraints from those clusters are based on their masses, dominated by dark matter. This gives rise to complications both in the estimation of cluster masses (which must be inferred from the observed baryon properties) and in the determination of sample selection functions. Incorporating selection functions in any cosmological model depends crucially on the form and scatter of the relationship between
the observable used to detect the cluster, and the cluster mass.

In principle, these complications may be avoided by defining cluster samples through gravitational lensing, the most direct observational probe of cluster masses. The development of weak lensing (abbreviated WL throughout) techniques has enabled the detection of clusters via the distortions of background galaxy shapes leading to the construction of WL selected cluster surveys. These include the Deep Lens Survey (Miyazaki et al. 2007), both \( \approx \) 1000 deg\(^2\), and the large (\( \sim 170 \) deg\(^2\)) CFHT Legacy Survey (Shan et al. 2012). However, while WL techniques are free from assumptions about the relationship between baryonic and dark matter in clusters, they are susceptible to the possibility of projection of multiple structures along the line of sight, leading to overestimates of cluster masses or false detections.

This paper aims to determine the X-ray properties of clusters detected via their WL signal. By constructing scaling relations based on the measured X-ray properties, we will determine whether the clusters follow simple scaling theory, which is a key ingredient in the determination of cosmological parameters. One of the most important measurements for use in cosmological studies is that of the cluster mass. Much work has been done to determine cluster masses from weak lensing observations, and the comparison to X-ray observables and X-ray determined masses, however this is primarily based on X-ray selected cluster samples (e.g. Hoekstra 2007; Okabe et al. 2010; Jee et al. 2011; Israel et al. 2012; Hoekstra et al. 2012; Mahdavi et al. 2013; Israel et al. 2014). We will compare the scaling relation between the X-ray luminosity and WL mass from WL selected and X-ray selected samples. The derivation of the WL cluster mass assumes that the clusters follow spherical symmetry. We investigate this assumption and determine the dynamical state and cool-core fraction of WL selected clusters.

The outline of this paper is as follows. In §2, we discuss the sample selection and data reduction. The derivation of the cluster properties is given in §3. In §4, we present the results of our X-ray analysis and derive scaling relations. Our discussions and conclusions are presented in §5 and §6 respectively. Throughout this paper we assume a cosmology with \( \Omega_M=0.3 \), \( \Omega_{\Lambda}=0.7 \) and \( H_0=70 \) km s\(^{-1}\) Mpc\(^{-1}\).

2 SAMPLE AND DATA REDUCTION

Our clusters were selected from the Subaru WL survey of Miyazaki et al. (2007). The sample defined in this work was constructed based upon confirmed cluster identification from spectroscopic follow up (Hamana et al. 2009). For 36 WL cluster candidates 15-32 galaxy redshifts were obtained per cluster, with 28 candidates securely identified as clusters. 10 of these clusters (within the redshift range \( 0.14<z<0.42 \)) were defined as a “clean” subset of clusters whose velocity dispersion could be evaluated from at least 12 spectroscopic member galaxies, and whose WL mass estimates are not affected by a neighboring system or field boundary. A summary of the 10 clusters investigated in this paper is given in Table 1. For eight of the clusters in the sample we obtained dedicated Chandra pointings, whereas SLJ1000.7 was observed in the COSMOS field with Chandra and SLJ0225.7 was observed in the XMM_LSS_13 field with XMM. We note however that the Chandra of SLJ1647.7 produced low quality data, we therefore chose to use an archived XMM observation for the analysis of this cluster.

For the eight clusters analysed using Chandra observations we used the CIAO\(^1\) 4.4.7 and followed standard reduction methods. Since all observations were telemetered in VFAINT mode additional background screening was applied\(^2\). We inspected background light curves of the observations following the recommendations given in Markevitch et al. (2003), to search for possible background fluctuations. None of the observations were affected by periods of background flaring.

In order to take the background into account, we follow the method described in Vikhlinin et al. (2006). Briefly, Chandra blank-sky backgrounds were obtained, processed identically to the cluster, and reprojected onto the sky to match the cluster observation. We then renormalise the background in the 9.5–12 keV band, where the Chandra effective area is nearly zero and the observed flux is due entirely to the particle background events. Finally, to take into account differing contributions from the soft X-ray background, the spectra were subtracted and residuals were modeled in the 0.4–1keV band using an unabsorbed APEC thermal plasma model (Smith et al. 2001) with a temperature of 0.18 keV. This model was included in the spectral fitting for the cluster analysis.

The clusters SLJ0225.7 and SLJ1647 were observed by XMM. We used the Science Analysis Software (SAS) version 12.0.1 and the most recent calibration products as of Oct. 2013 for the analysis of these clusters.

Table 1 lists the total cleaned exposure times for each cluster.

3 CLUSTER PROPERTIES

In this section we outline the methods used to determine the cluster properties of our sample. Where possible we use the method outlined in Section 3.1, however for clusters with low SNR data, the properties are determined using one of the methods described in Section 3.2. Table 1 lists the analysis method used for each cluster.

3.1 The “Standard” Method

To determine the cluster properties of our sample we followed the procedures outlined in Maughan et al. (2012), for the Chandra observations. We detail our analysis for the XMM observations of the clusters SLJ0225.7 and SLJ1647.7 in Sect. 3.3.1 and 3.3.0 respectively. Briefly, cluster spectra were extracted and fits performed in the 0.6 - 9.0 keV band with an absorbed APEC model, with the absorbing column fixed at the Galactic value (Kalberla et al. 2005). We note that due to the low SNR of five of the clusters (see 3.3), the abundance was fixed at 0.3\(^\oplus\) throughout this analysis. We determine the gas density profile for each cluster by converting the observed surface brightness profile (constructed

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1 See http://cxc.harvard.edu/ciao/
2 See http://cxc.harvard.edu/caldb/
3 See http://cxc.harvard.edu/ciao/why/aciscleanv1.html

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in the 0.7 – 2.0 keV band) into a projected emissivity profile, which is modeled by projecting a density model along the line of sight. We use the model of Vikhlinin et al. (2006) see that work for parameter definitions):

\[ r_p n_0 = \frac{r_50^2 (r/r_50)^{-\alpha}}{(1 + r^2/r_50^2)^{3\beta - \alpha/2}} \times (1 + r^2/r_50^2)^{-\gamma} \]  

The same parameter constraints were employed as in Vikhlinin et al. (2006) i.e. \( \gamma \) is fixed at 3 and \( \epsilon < 5 \) to exclude nonphysical sharp density peaks. Gas masses were determined from Monte Carlo realisations of the projected emissivity profile based on the best fitting projected model to the original data.

The cluster temperature, gas mass and \( r_{500} \) (the radius at which the enclosed density of the cluster becomes 500 times the critical density at the cluster’s redshift) were determined through an iterative process. We extract a spectrum from within an estimate of \( r_{500} \) (with the central 15% excluded), integrate a gas density profile (see Maughan et al. 2008) to determine the gas mass, and then calculate a value for \( Y_X \) (the product of the temperature and gas mass; Kravtsov, Vikhlinin & Nagai 2006). A new \( r_{500} \) was then estimated from the \( Y_X \)–M scaling relation of Vikhlinin et al. (2009a).

\[ M_{500} = E(z)^{-2/5} A_Y M \left( \frac{Y_X}{3 \times 10^{44} M_\odot \text{keV}} \right)^{B_Y M} \]  

with \( A_Y = 5.77 \times 10^{14} M_\odot \) and \( B_Y M = 0.57 \). Here, \( M_{500} \) is the mass within \( r_{500} \), where \( r_{500} \) is defined as \( r_{500} = (3M_{500}/4\pi 500 \rho_c(z))^1/3 \), where \( \rho_c(z) \) is the critical density of the universe at the cluster redshift. The \( Y_X \)–M relation assumes self-similar evolution (corrected by \( E(z)^{-2/5} \)), where \( E(z) = \sqrt{\Omega_M (1+z)^3 + (1-\Omega_M-\Omega_\Lambda)(1+z)^2 + \Omega_\Lambda} \) (justified by Maughan 2007).

An initial temperature of 2 keV was used to determine an initial \( r_{500} \) and then this process was repeated until \( r_{500} \) converged to within 1%. The luminosity and temperature were measured from spectra extracted within \( r_{500} \) both with and without the central 15% of \( r_{500} \) excluded. Throughout we define \( L_X \) and \( kT \) as core excluded cluster properties, and \( L_X Z \) and \( kT \) as core included properties. All luminosities quoted throughout are bolometric unless otherwise stated.

### 3.2 The “Low SNR” Method

Due to the low SNR of the clusters resulting from the clusters being less luminous than expected based on the WL mass (used when planning the observations), five of the clusters in our sample could not be analysed using the method described in Sect. 3.1. We therefore used a variety of methods in order to obtain the most accurate properties for these low SNR clusters. In order of preference, these were:

(a) When the gas mass could not be reliably measured, we used the temperature alone to measure \( r_{500} \). We used the \( r_{500} \)–\( T \) relation given in Vikhlinin et al. (2006), and followed the iterative process described in Sect. 3.1 to determine the cluster properties within \( r_{500} \). We note that for consistency with the \( r_{500} \)–\( T \) relation, the central 70 kpc is excluded (instead of 0.15%\( r_{500} \) as in § 3.1).

(b) Extracting a spectrum within the highest SNR region for the cluster and determining the properties within this region. \( r_{500} \) was then estimated using the \( r_{500} \)–\( T \) of Vikhlinin et al. (2006) and \( L_X \) was extrapolated to the radius by integrating under a \( \beta \)-profile with \( r_c = 150 \text{ kpc} \) and \( \beta = 0.667 \). Again, the central 70 kpc is excluded for this analysis.

(c) Using PIMMS\(^4\) to determine the cluster luminosity from the count rate of the cluster observation when no spectroscopic analysis could be performed. A global temperature of 2 keV was assumed for the cluster, and the count rate determined in an \( r_{500} \) determined from the \( r_{500} \)–\( T \) relation of Vikhlinin et al. (2006).

### 3.3 Notes on Individual Clusters

Notes on the WL detections of the individual clusters can be found in Hamana et al. (2009). In this section we note any peculiarities or points of interest for observations in which we departed from analysis process described in Sect. 3.1. We note that five of the observations in our cluster sample were analysed using the method described in Sect. 3.2 and these clusters are discussed below.

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\(^4\) http://cxc.harvard.edu/toolkit/pimms.jsp
3.3.1 SLJ0225.7–0312

The cluster SLJ0225.7–0312 had an archival \textit{XMM} observation taken in the XMM\textsubscript{LSS}13 field. An image of the \textit{XMM} observation is shown in Figure 3\textsubscript{a}. The properties of the cluster were obtained following method (b) outlined in Sect. 3.2. A spectrum was extracted within [0-433]kpc for each of the \textit{XMM} cameras and the spectra fit simultaneously with the temperature of the hot APEC components tied together. The net counts in the 0.3–7.0 keV band were 2635, 1019 and 1053 for the pn, mos1 (m1) and mos2 (m2) cameras respectively, with SNR 40, 24 and 25 for the pn, m1 and m2 cameras respectively. Due to the large field of view of the \textit{XMM} cameras, a local background region was used for background subtraction. We found a temperature of \(kT=2.93\pm0.19\) keV for the cluster, corresponding to an \(r_{500}\) of 798 kpc. We found an extrapolated luminosity within \(r_{500}\) for the cluster of \(L_X=(7.31\pm0.19)\times10^{45}\) ergs s\(^{-1}\). The analysis carried out above was also employed to determine the properties of the cluster with the core included. We derived a temperature of \(kT_c=3.00^{+0.18}_{-0.17}\) keV. Extrapolating outwards to \(r_{500}\), we determined a luminosity of \(L_{X,c,}(7.90\pm0.19)\times10^{45}\) ergs s\(^{-1}\).

3.3.2 SLJ1000.7+0137

The cluster SLJ1000.7+0137 had three archival \textit{Chandra} observations taken in the COSMOS\textsubscript{6} field. An image of the \textit{Chandra} observation is shown in Figure 3\textsubscript{a}. The three individual observations were analysed separately as described in Sect. 3.2. The data were then combined for certain stages of the analysis. Source, background and exposure maps were projected onto common coordinate systems and summed. Source and background spectra were extracted for individual observations and fit simultaneously with the temperatures of the hot APEC components tied together.

3.3.3 SLJ1135.6+3009

In order to calculate the luminosity of this cluster, we followed method (c) outlined in Sect. 3.2. We determined the number of counts in a region centered on the location of the Brightest Central Galaxy (BCG, see Figure 3\textsubscript{a}) with the BCG highlighted by the black circle), subtracting the number of counts from the same region in a blank-sky background obtained for the observation, scaled by the ratio of the source and background exposure times. The net number of counts obtained was \(71\pm22\) (significant at the 3.2\textsigma level), in the 0.7–2.0 keV band. We calculated a bolometric luminosity of \(L_{X,c}=(1.28\pm0.40)\times10^{43}\) ergs s\(^{-1}\) for the cluster, with the errors derived assuming Poisson statistics.

3.3.4 SLJ1337.7+3800

To determine the cluster properties we used method (b) in Sect. 3.2 extracting spectra within a radius [70-215]kpc. The net number of spectral counts, corrected for background, was then 130 (in the 0.6–9.0 keV band, with SNR=5.3). We find a temperature of \(kT=1.63^{+0.06}_{-0.02}\) keV, corresponding to an \(r_{500}\) of 574 kpc. We derive an extrapolated bolometric luminosity of the cluster of \(L_X=(6.58\pm1.31)\times10^{42}\) ergs s\(^{-1}\). This analysis was repeated including the core (obtaining 172 net spectral counts in the 0.6–9.0 keV band, with SNR=6.45). We determined a temperature of \(kT_c=1.63^{+0.47}_{-0.23}\) keV and an extrapolated bolometric luminosity of \(L_{X,c}=(8.56\pm1.21)\times10^{42}\) ergs s\(^{-1}\).

3.3.5 SLJ1634.1+5639

Due to observational constraints, the exposure was split into two observations. Therefore the analysis of the cluster was performed as described in Sect. 3.3.2. The properties of this cluster were obtained following the process outlined in method (a) in Sect. 3.2. We find a temperature of \(kT=1.37^{+0.40}_{-0.44}\) keV and \(L_X=(4.17\pm1.67)\times10^{42}\) ergs s\(^{-1}\). Including the core of the cluster we found a temperature and luminosity of \(kT_c=1.72^{+0.42}_{-0.93}\) keV and \(L_{X,c}=(6.78\pm2.59)\times10^{42}\) ergs s\(^{-1}\).

Substructure in SLJ1634 appears likely on the basis of the locations of the galaxies, as marked in Figure 4\textsubscript{a} which shows a split in the population of galaxies in the northern and southern region of the cluster. To check for the reality (or otherwise) of this structure, a tree analysis of the locations of the galaxies in 3D (position, velocity, redshift) space was undertaken. The results suggest a composite structure for the cluster, with the galaxies separated into three separate groups. The galaxies in Figure 4\textsubscript{a} are colour coded with respect to each group they are associated. Group A (magenta squares), B (black squares) and C (green squares) have a redshift of 0.231, 0.238 and 0.242 respectively. This composite structure may explain the large offset in the luminosity-mass relation for SLJ1634 when the WL mass is calculated centered on the X-ray peak (see Fig. 6).

3.3.6 SLJ1647.7+3455

Although a dedicated \textit{Chandra} observation was obtained for the cluster SLJ1647.7+3455, due to the non-detection of the cluster, no useful information could be determined. We therefore used an archival \textit{XMM} observation to determine the properties of the cluster, following the method outlined in Sect. 3.2. The spectra were analysed in the same way as in Sect. 3.3.1. We found the temperature of \(kT=1.49^{+0.11}_{-0.11}\) keV and \(L_X=(1.38\pm0.12)\times10^{42}\) ergs s\(^{-1}\). Including the cluster core, we find \(kT_c=1.58^{+0.09}_{-0.09}\) keV and \(L_{X,c}=(1.57\pm0.12)\times10^{42}\) ergs s\(^{-1}\).

4 RESULTS

In this section we present the results of our analysis of our 10 WL selected clusters. We derive various scaling relations for these clusters and compare to published results. The measured properties of the clusters are given in Table 2 along with the WL mass determined in Hamana et al. (2009), converted to \(H_0=70\) km s\(^{-1}\) Mpc\(^{-1}\). We note two updates to the derivation of the WL masses given in Hamana et al. (2009), and those presented in this work. We use the halo mass-concentration relation of Duffy et al. (2008), and the modified NFW profile given in Oguri & Hamana (2011).
Table 2. Derived X-ray properties of the clusters with WL mass estimates (see §4). Clusters in the top part of the table were analysed from XMM observations, in the bottom part from Chandra observations.

<table>
<thead>
<tr>
<th>Cluster</th>
<th>E(z)</th>
<th>$L_X$ ($\times 10^{43}$ ergs s$^{-1}$)</th>
<th>$L_{X,c}$ ($\times 10^{43}$ ergs s$^{-1}$)</th>
<th>$L_{0.1-2.4keV}$,$X_{c}$ ($\times 10^{43}$ ergs s$^{-1}$)</th>
<th>$kT$ (keV)</th>
<th>$kT_c$ (keV)</th>
<th>$M_{WL,500}$ ($\times 10^{14}$M$_\odot$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>SLJ0225.7-0312</td>
<td>1.070</td>
<td>7.31±0.19</td>
<td>7.90±0.19</td>
<td>4.16±0.10</td>
<td>2.93±0.19</td>
<td>3.00±0.18</td>
<td>1.99±0.47</td>
</tr>
<tr>
<td>SLJ1647.7+3455</td>
<td>1.140</td>
<td>1.58±0.12</td>
<td>1.57±0.12</td>
<td>1.31±0.09</td>
<td>1.40±0.11</td>
<td>1.58±0.10</td>
<td>2.00±0.79</td>
</tr>
<tr>
<td>SLJ0850.5+4512</td>
<td>1.100</td>
<td>0.78±0.14</td>
<td>0.92±0.13</td>
<td>0.83±0.13</td>
<td>1.12±0.22</td>
<td>1.16±0.14</td>
<td>1.09±0.39</td>
</tr>
<tr>
<td>SLJ1000.7+0137</td>
<td>1.114</td>
<td>4.04±0.17</td>
<td>5.62±0.19</td>
<td>3.80±0.17</td>
<td>2.43±0.29</td>
<td>2.61±0.41</td>
<td>2.39±0.46</td>
</tr>
<tr>
<td>SLJ1135.6+3009</td>
<td>1.108</td>
<td>–</td>
<td>1.28±0.40</td>
<td>1.08±0.26</td>
<td>–</td>
<td>–</td>
<td>2.49±0.56</td>
</tr>
<tr>
<td>SLJ1204.4-0351</td>
<td>1.141</td>
<td>3.43±0.35</td>
<td>3.97±0.38</td>
<td>2.92±0.28</td>
<td>2.16±0.39</td>
<td>2.16±0.59</td>
<td>1.20±0.60</td>
</tr>
<tr>
<td>SLJ1335.7+3731</td>
<td>1.239</td>
<td>3.10±0.71</td>
<td>3.99±0.69</td>
<td>3.35±0.68</td>
<td>1.39±0.63</td>
<td>1.78±0.47</td>
<td>2.79±1.01</td>
</tr>
<tr>
<td>SLJ1337.7+3800</td>
<td>1.092</td>
<td>0.66±0.13</td>
<td>0.86±0.12</td>
<td>0.37±0.05</td>
<td>1.63±0.22</td>
<td>1.63±0.20</td>
<td>1.24±0.30</td>
</tr>
<tr>
<td>SLJ1602.8+4335</td>
<td>1.245</td>
<td>12.7±1.14</td>
<td>16.4±1.22</td>
<td>8.50±0.54</td>
<td>3.91±0.59</td>
<td>4.45±1.43</td>
<td>2.60±0.53</td>
</tr>
<tr>
<td>SLJ1634.1+5639</td>
<td>1.126</td>
<td>0.63±0.22</td>
<td>0.68±0.26</td>
<td>0.49±0.15</td>
<td>1.37±0.23</td>
<td>1.42±0.23</td>
<td>0.87±0.49</td>
</tr>
</tbody>
</table>

4.1 The Luminosity-Temperature Relation

Here we derive the X-ray luminosity-temperature relation for our clusters, omitting SLJ1135 for which a temperature could not be obtained. Figure 1 shows the $L_{X,c}$ vs $kT_c$ relation for our clusters (red triangles) compared with the LT relation given in Maughan et al. (2012 blue open circles). The Maughan et al. (2012) sample of clusters was selected from all available Chandra pointings as of 2006 November, correlated with the NASA/IPAC Extragalactic Database (NED) and observations with a galaxy group or cluster listed in NED within 10’ of the Chandra aimpoint were kept. Further criteria included observations only carried out with the ACIS-I detector, a galaxy group/cluster as the target, and a redshift criteria included observations only carried out with the ACIS-I detector, a galaxy group/cluster as the target, and a redshift $>0.1$. As the Maughan et al. (2012) sample of clusters covers a wide redshift range (0.1 < z < 1.3), the luminosities of the clusters were corrected for the expected self-similar evolution, given by $L_X \propto E(z)^{-1}$. The same correction was also applied to our data.

A power law of the form $L_X = E(z)^2 L_0 (kT/kT_0)^{B_{LT}}$ was fit to the data using the BCES orthogonal regression in log space (Akritas & Bershady 1996) with $kT_0$ set at 2 keV. From our fit we find a normalisation of $L_0=(2.44\pm0.51)\times10^{43}$ ergs s$^{-1}$ and a slope of $B_{LT}=2.63\pm0.69$. The slope and normalisation are consistent with the fit to the Maughan et al. (2012) sample, who find a normalisation of $L_{0,M12}=(2.40\pm0.78)\times10^{43}$ ergs s$^{-1}$ at 2 keV and a slope of $B_{LT,M12}=3.63\pm0.27$ (see Fig. 2 blue solid line). We conclude that the X-ray properties of the WL selected clusters scale consistently with X-ray selected clusters in the luminosity-temperature plane.

4.2 The Luminosity-Mass Relation

Here we derive the luminosity-mass (LM) relation for our clusters, with the masses estimated from the WL analysis presented in Hamana et al. (2009), updated as described in Section 4. Figure 2 shows the luminosity-mass relation for our clusters (red triangles), and compared to the data given in Mantz et al. (2010a black crosses). Their clusters are taken from the ROSAT Brightest Cluster Sample (BCS; Ebeling et al. 1998), the ROSAT-ESO Flux-Limited X-ray Sample (REFLEX; Böhringer et al. 2004), and the bright sub-sample of the Massive Cluster Survey (Bright MACS; Ebeling et al. 2010), all $z<0.5$. We note that Mantz et al. (2010a) uses the gas mass as a robust, low-scatter proxy for the total cluster mass. The luminosities in Mantz et al. (2010a) are calculated in the 0.1–2.4 keV band (cluster rest frame). We therefore calculate the luminosities of the WL clusters in this band for the purpose of this comparison.

Since we wish to compare weak-lensing and X-ray selected clusters, we need to minimise any selection effects in the X-ray comparison sample. X-ray selected cluster samples suffer from both Malmquist and Eddington biases. Malmquist bias refers to the fact that a larger survey volume finds more luminous objects for a given flux limit, and Eddington bias refers to the differential scattering of objects across the detection threshold of a survey due to the slope of the mass function. The analysis presented in Mantz et al. (2010a) takes fully into account these biases, making it an ideal comparison for this work. We also note that the correction for the E(z) factor in Mantz et al. (2010a) for the LM relation was based upon the best fit cosmology found in
Mantz et al. (2010b). For the purposes of the LM relation comparison, the $E(z)$ factor for each cluster is also calculated using the Mantz et al. (2010b) cosmology. On average, $E(z)$ changes by $\approx 3\%$.

A power law of the form $E(z)^{-1}L_X = L_0(E(z)/M/M_\odot)^{B_{LM}}$ was fit to the WL clusters in log space, with $L_0$ set at $2\times10^{44}M_\odot$. Due to the low quality of the data and the small number of data points we fixed the slope ($B_{LM}$) of the relation at the self similar expectation of 1.33. We note that since the slope is fixed, we perform a $\chi^2$ fit to the cluster LM relation. For a mass of $2\times10^{14}M_\odot$ we found $L_0=(1.9\pm0.2)\times10^{43}$ ergs s$^{-1}$ (see red solid line Figure 2). The fit from Mantz et al. (2010b) is represented by the solid black line. The low normalisation compared to the data demonstrates the size of the Mahnquist and Eddington biases (which were included in the analysis of the Mantz et al. (2010b) data. The Mantz et al. (2010b) LM relation at the same mass predicts $L_0=7.3\times10^{43}$ ergs s$^{-1}$. Compared with X-ray selected clusters, X-ray luminosity is lower at a given mass by a factor 3.9$\pm$0.9. Conversely, $M_{WL,500}$ is higher for a given luminosity by a factor 2.9$\pm$0.2. We discuss the nature of this large offset in the LM relation in Sect. 5.

The main drawback in the comparison with the Mantz et al. (2010b) sample of clusters is the mass range covered by the respective cluster samples. The Mantz et al. (2010b) sample of clusters have masses $M_{500} \geq 2.7\times10^{14}M_\odot$, whereas our WL sample are all $<2.7\times10^{14}M_\odot$. To ensure the mass range of the comparison sample is not playing a role in the large observed offset, we compare to a sample of clusters probing the lower mass end and a sample of galaxy groups. We use the REXCESS (Bohringer et al. 2007) sample of clusters studied in Pratt et al. (2009) which span the mass range $\approx 10^{14}-10^{15}M_\odot$, and the sample of galaxy groups presented in Eckmiller, Hudson & Reiprich (2011) spanning the mass range $\approx 0.05-2\times10^{14}M_\odot$. The masses of the Pratt et al. (2009) clusters were calculated using the $Y_X$-Mass relation given in Arnaud, Pointecouteau & Pratt (2007), and the masses of the Eckmiller, Hudson & Reiprich (2011) groups were calculated by assuming hydrostatic equilibrium and spherical symmetry (see Sect. 3.4 in Eckmiller, Hudson & Reiprich 2011). In both cases the luminosities are calculated in the 0.1-2.4 keV band (cluster rest frame). We plot the clusters on the LM relation in Figure 2 with the Pratt et al. (2009) clusters given by the green squares and the Eckmiller, Hudson & Reiprich (2011) groups given by the blue stars. A visual inspection shows that all the X-ray selected comparison groups lie on approximately the same LM relation. Although the Pratt et al. (2009) and Eckmiller, Hudson & Reiprich (2011) samples appear to lie slightly above the Mantz et al. (2010b) relation, this is consistent with a similar amount of bias as in the Mantz et al. (2010b) sample, with the relation correcting for these effects. The offset of

![Figure 2](https://example.com/figure2.png)
the WL selected clusters is still evident down to these low masses.

4.3 X-ray Morphology of Weak Lensing Selected Clusters

4.3.1 The dynamical state of weak lensing selected clusters

We wish to investigate the dynamical state of our WL selected clusters as compared to X-ray selected cluster samples. We use the centroid shift method of Poole et al. (2006) to determine the dynamical state of the cluster. The centroid shift \((w)\) was defined as the standard deviation of the distance between the X-ray peak and centroid, measured within a series of circular apertures centered on the X-ray peak and decreasing in steps of 5% from \(r_{500}\) to 0.05\(r_{500}\). The \((w)\) values are given in Table 3. Maughan et al. (2012) used \((w)\) to distinguish between “relaxed” and “unrelaxed” clusters. For the Maughan et al. (2012) sample of 114 clusters they adopted a value of \((w)=0.006r_{500}\) below which clusters are classed as “relaxed”, and above which are classed as “unrelaxed”. Using this threshold value, 25% of the Maughan et al. (2012) sample was classed as “relaxed”. If we adopt this value for our sample of WL selected clusters, we find that none of the clusters in Table 3 qualify as “relaxed”.

The value of \((w)\) is somewhat subjective, and the value used to separate relaxed and unrelaxed clusters can depend on the sample of clusters used, and indeed on each individual study. Pratt et al. (2009) used the REXCESS sample of clusters and defined relaxed/unrelaxed clusters using the same method of calculating \((w)\) as we employ here. Pratt et al. (2009) found that ~60% of their clusters has \((w)<0.01\), while for our WL clusters we find that only one of the seven clusters with measurable \((w)\) values are below this threshold. We note however that this comparison is only based on a small number of clusters from the WL.

This result appears analogous to the recent cluster sample from the early release Planck all sky survey. Planck Collaboration et al. (2011b). These clusters were detected via their SZ signals, and 25 candidates were followed up with snapshot XMM observations. Through a visual inspection of the gas density profiles constructed for each cluster, a large proportion of these SZ selected clusters appear morphologically disturbed. Since the WL selected clusters studied in our sample also appear to be morphologically disturbed, selecting clusters from either their SZ or WL signal, as compared to selecting cluster via their X-ray emission, appears to give a more representative distribution of morphologies for clusters.

4.3.2 The cool-core fraction of weak lensing selected clusters

The absence of relaxed clusters in our sample as indicated by \((w)\) suggests we should also see few cool core (CC) clusters. This also has important implications for cosmology as the CC fraction and its evolution may be over-represented/biased in X-ray samples (Santos et al. 2010). This result is not unexpected due to the high X-ray surface brightness at the center of CC clusters, making them easier to detect in X-ray cluster surveys. Hudson et al. (2010) presented a comprehensive study of 16 CC diagnostics for the HIFLUGCS sample of 64 clusters. They found that, for clusters with low data quality, the cusplines of the gas density profile is the preferred method of determine the presence of a CC. Cuspliness is defined as the logarithmic slope of the gas density profile at a radius of 0.04\(r_{500}\), modeled using the gas density profiles given by equation 4 with the uncertainties derived from the cusplines values measured from Monte Carlo realisations of the gas density profile. We obtained the cusplines values from gas density profiles fitted to surface brightness profiles derived from images binned by a factor 2, with each radial bin containing at least 10 cluster counts. The same fitting process was followed as in 5.1 and the cusplines values were derived from these profiles. Table 3 lists the values of the cusplines for each cluster where a gas density profile could be obtained.

Maughan et al. (2012) also used the cusplines to determine the presence of a CC in their sample of 114 clusters. Above a value of 0.8 clusters were considered to host a CC, and below they were not. If we adopt this value for our sample of 6 clusters with measurable cusplines, we find that one of the clusters is considered to host a CC. We note that the Maughan et al. (2012) sample spans a wide range in redshift and data quality, and again the threshold in cusplines used to define CC/NCC clusters is subjective depending on the sample being studied. The Hudson et al. (2010) sample consists of clusters with high data quality and the value of the cusplines is well constrained for all clusters. Hudson et al. (2010) adopted a value for cusplines of 0.7 to segregate between CC and non-CC clusters. If we adopt this value of cusplines for our sample of clusters, we find only 2/6 of the clusters host a CC. In both cases we find a smaller percentage of CC clusters in comparison to these samples of clusters, we find 16.6% of clusters host a CC when using a cusplines value of 0.8, as compared to 26.3% found in Maughan et al. (2012). We find that 33.3% of clusters host a CC when using a cusplines value of 0.7, as compared to 54.7% found in Hudson et al. (2010). This suggests that CC clusters may be over-represented in X-ray selected samples of clusters.

5 DISCUSSION

In Sect. 1.2 we showed that WL selected galaxy clusters are underluminous by a factor \(\approx 3.9\) for a given mass, compared to X-ray selected clusters. In terms of the mass, we find
that the WL selected clusters are over-massive by a factor \(\approx 3\) compared to X-ray selected clusters. The consistency of the \(L_{X,c}-kT_c\) relation with X-ray selected clusters (see Fig. 1) suggests that the discrepancy may more likely arise as a result of a bias in the mass. We therefore first investigate the nature of this offset through a study of systematic effects on the WL mass estimates.

5.1 Possible Systematics in the WL Masses

In the following sections we present a series of plausible effects that could bias the WL mass. These biases are presented as a cumulative effect on the WL mass, and we investigate the effect these biases have on the LM relation.

5.1.1 The Effect of Cluster Centre Position on Weak Lensing Mass

The X-ray analysis presented in this paper is independent of the WL analysis given in Hamana et al. (2009), and as such, different locations of the cluster centre are used for the two analysis. The WL masses given in Table 2 are derived from fitting NFW models to shear profiles centered on the peak position in the WL mass maps, whereas the luminosities are derived centered on the X-ray centroid. Here we investigate the effect of deriving the WL masses from profiles centered on the X-ray centroid.

Figure 3 shows the ratio of the WL mass centered on the peak of the WL signal to the WL mass centered on the X-ray centroid \(\left(\frac{M_{500}^\text{WL peak}}{M_{500}^\text{Xcent}}\right)\) against the offset of the WL peak and X-ray centroid (in units of 500). The cluster SLJ1634 is highlighted by the black circle. This cluster shows evidence for a composite structure (see § 3.3.5) and the WL mass decreases by a factor \(~7\) when centered at the X-ray peak. We note for the cluster SLJ1135, the X-ray analysis was centered on the WL peak due to the low SNR of the X-ray data, and no X-ray centroid could be obtained. It is thought that any offset between the WL and X-ray would be minimal since the WL peak is aligned closely with the BCG of the cluster (see Fig. A1(b)). This figure shows that the ratio of the masses is consistent with a ratio of 1.0 (given by the black dashed line). This result is unsurprising due to the radial range within which the WL masses are derived (2\(<\theta<20^\circ\)), to neglect contamination from cluster galaxies. These results are consistent with the results from Okabe et al. (2010a), who find that the relative inaccurate of cluster centroid position has a negligible effect on the resulting \(\chi^2\) value of the NFW fit to their distortion profiles, and hence the mass derivation.

The effect of using WL masses centered on the X-ray centroid on the LM relation is shown in Figure 3 as the dotted blue line. The original LM relation from Figure 2 is given by the dashed green line. By calculating the WL masses centered on the X-ray centroids, the masses have decreased on average by a factor of \(~1.1\).

5.1.2 "Deboosting" the Weak Lensing Mass Estimates

Here we correct the weak lensing masses by "deboosting" the shear signal, which is analogous to the effect of Eddington bias (see Sect 4.2). Since our sample is based upon a noisy indicator of the cluster mass, the shear signal \((\kappa)\), clusters with high mass but a low observed \(\kappa\) will not be included in our sample, and conversely clusters with low masses but high observed shear will be included in our sample. Our cluster sample will therefore be biased to clusters with a high shear signal relative to their mass, and due to the slope of the mass function, there will be more low mass than high mass clusters close to our detection limit. The overall effect will be to measure high mass values, due to the high shear signal, and so bias the sample to higher masses. An application of "deboosting" can be found in Coppin et al. (2004) for a SCUBA source catalog.

In order to calculate "deboosted" mass estimates, we follow a similar method to that carried out in Coppin et al. (2006). Initially, the cluster mass is estimated from the shear profile \((\kappa)\) via \(\chi^2\) fitting with an NFW model. We also have a theoretical prediction of the shear profile for a modified NFW profile. We compute the \(\chi^2\) following the standard procedure from the observed shear profile and the model prediction as a function of halo mass \(M\) (\(\chi^M\)). From \(\chi^M\) we compute the best-fit (non-"deboosted") WL mass and likelihood interval. Next we consider \(P = \exp(-\chi^2/2)\) as \(P(\kappa, \sigma | M)\), and consider the halo mass function as \(p(M)\). From Bayes’ theorem we therefore have \(P(M | \kappa, \sigma) = P(M)P(\kappa, \sigma | M)\). We then compute \(\chi_{\text{deboost}}^2 = -2\ln(P(M | \kappa, \sigma))\) and find the best fit mass and the likelihood function.

This correction decreases the WL mass by an average of \(~1.27\). This is shown in Figure 5 as the dotted pink line, representing the cumulative effect of the centroid position (see § 5.1.1) and Eddington bias on the LM relation.

5.1.3 Weak Lensing Biases due to Halos Elongated Along the Line of Sight

An additional bias involved in the estimation of weak lensing masses is that due to a triaxial shape of the cold dark matter (CDM) halo (Warren et al. 1992; Jing & Suto 2002; Shaw et al. 2006). This can lead to errors in the weak lensing masses by \((10-50)\%\) when spherical symmetry is assumed.
As stated in §5.1.4, the original WL analysis was independent of the X-ray analysis. The LM relation shown in Figure 2 had the masses derived within a WL derived $r_{500}$ ($r_{500}^{WL}$) and the luminosities calculated within an X-ray derived $r_{500}$ ($r_{500}^{X}$). In order to compare consistently with the Mantz et al. (2010b) LM relation, whose masses were derived within $r_{500}$, we recalculate the WL masses within $r_{500}$ for each cluster. Figure 4 plots the change in WL mass against the ratio of the WL and X-ray derived $r_{500}$. The WL masses decrease on average by a factor of $\approx 1.6$. The correction for centroid position, Eddington bias, triaxiality and the calculation of the WL mass within $r_{500}$, is given by the red solid line in Figure 5.

From Figure 5 we can clearly see the largest correction to the WL masses is when they are determined within X-ray derived $r_{500}$ values. However, this partially obscures the true discrepancy. This shows that the agreement between WL and X-ray masses is improved if the same aperture is used for both, but this is no longer a WL $M_{500}$. We investigate this further in Section 5.3.

5.1.5 Corrected Luminosity-Mass Relation

In Sections 5.1.1 and 5.1.4 we presented a series of systematic and statistical corrections to our WL cluster masses. We re-derive the LM relation for the WL selected clusters, with the WL masses corrected for these effects. Figure 5 plots the corrected LM relation, with each of the lines corresponding to a fit with the cumulative effect of these corrections taken into account. From the final fit to the mass corrected LM relation (Fig 5 red solid line), we find
Table 5. Observed scaling relations. Each scaling relation is fit with a power law relation of the form $E(z)\propto L_X = A_0(X/X_0)^\beta$, where $X_0$ is 2 keV and $2\times10^{14}$ $M_\odot$ for kT and $M_X$ respectively.

<table>
<thead>
<tr>
<th>Normalisation</th>
<th>Slope</th>
</tr>
</thead>
<tbody>
<tr>
<td>$L_{X,c} - kT_c$</td>
<td>$A_{LT_{X,c}}(10^{43})$</td>
</tr>
<tr>
<td>Centroid Corrected</td>
<td>$A_{LM}(10^{43})$</td>
</tr>
<tr>
<td>Centroid and Edd. Bias Corrected</td>
<td>$A_{LM}(10^{43})$</td>
</tr>
<tr>
<td>Centroid, Edd. Bias and Ellipticity Corrected</td>
<td>$A_{LM,corr}(10^{43})$</td>
</tr>
</tbody>
</table>

Figure 5. Figure showing the LM relation for the WL selected clusters (red triangles), with the WL masses corrected for centroid position (§ 5.1.1), Eddington bias (§ 5.1.2) and triaxiality (§ 5.1.3). Finally, the WL masses are derived within the X-ray determined $r_{500}$ (see §5.1.4). The cluster SL1634 is highlighted by the black circle. A power law fit for the LM relation with the masses corrected for these effects is given by the red solid line. Once again compared to the clusters given in Mantz et al. (2010b) (black crosses), and corresponding bias corrected LM relation (black line). The clusters from Pratt et al. (2009) and groups from Eckmiller, Hudson & Reiprich (2011) are also plotted, but shaded grey for clarity.

$A_{LM,corr} = (4.7\pm0.6)\times10^{43}$ ergs s$^{-1}$, with the slope fixed at $\beta=1.33$. The comparison with the Mantz et al. (2010b) LM relation has been greatly reduced, and we find that the WL selected clusters are now consistent with X-ray selected clusters samples, albeit with three clusters being low luminosity outliers. We note that the clusters are only consistent when using WL masses derived within $r_{500}^{x}$, which may obscure the underlying discrepancy. These corrections therefore show that clusters detected from WL surveys suffer from biases which have a large effect on the derived LM relation. By considering a series of plausible biases on the WL mass, the discrepancy in the LM relation, when compared to X-ray selected clusters, can be resolved.

5.2 Further Potential Mass Biases

In the following sections we discuss further biases which may have an effect when comparing WL and X-ray hydrostatic mass calculations.

5.2.1 WL Mass Bias due to Large-Scale Structure

It has been shown that correlated and uncorrelated large-scale structure (LSS) along the LOS, can both induce bias and scatter in the WL mass (e.g. Metzler et al. 1999; Hoekstra 2001; White & Vale 2004; de Jager et al. 2005) Noh & Cohn 2011; Becker & Kravtsov 2011). We should also note that the effects of LSS and triaxiality are related, where it can be considered that the major axis of halos is correlated with filamentary LSS (e.g. Splinter et al. 1997; Faltenbacher et al. 2002; Hopkins, Bahcall & Bode 2005; Lee & Evrard 2007; Zhang et al. 2009; Noh & Cohn 2011). The distribution of BCG ellipticities suggests that 60% of our clusters are viewed with the major axis close to the LOS, from which we can assume that they will also be affected by LSS along the LOS. However, Meneghetti et al. (2010) note that since they consider all the mass within 20 $h^{-1}$ Mpc in their lensing simulations, the effect of correlated LSS is partially included in their estimate of the lensing mass. Therefore, the correction applied to our WL masses in Section 5.1.3 based upon Meneghetti et al. (2010), should partially include corrections for LSS. The effect of uncorrelated LSS has been found to not bias the WL mass, but adds additional scatter ≈15-30% as a function of cluster mass (Hoekstra et al. 2011; Becker & Kravtsov 2011).

5.2.2 Underestimates of X-ray Hydrostatic Mass

Simulations have shown that hydrostatic mass estimates of galaxy clusters are systematically biased towards underestimates of the true mass (e.g. Rasia, Tormen & Moscardini 2004; Kay et al. 2004; Nagai, Vikhlinin & Kravtsov 2007). Since hydrostatic estimates only take into account thermal pressure, the presence of additional pressure from random gas motions is neglected, and hence bias the mass low. It has been shown from simulations that up to ≈5%-20% of pressure support comes from random gas motions (e.g. Rasia et al. 2006; Lau, Kravtsov & Nagai 2009). However, the picture from observations is unclear. A recent study of 22 clusters detected by Planck, found that the average X-ray hydrostatic masses were 0.688±0.072 times lower than the masses determined from WL by von der Linden et al. (2014). Since the masses of the Mantz et al. (2010b) clusters are derived from a relation based upon X-ray hydrostatic mass estimates, in principle we can assume that their cluster masses are underestimated at the 10% level.

As stated in Section 4.2 the masses in Mantz et al.
of \( r_{500} \) had the largest effect on the derived WL mass and may obscure any further underlying biases. By making this comparison we can remove the large effect of the choice of \( r_{500} \). Furthermore, by comparing to WL masses the effects of hydrostatic mass bias (see §5.2.2), and to a lesser extent the WL mass bias due to LSS (see §5.2.1) but could still be larger in our sample compared to Mahdavi et al. (2013), due to the WL selection) will be reduced or removed.

Figure 6 plots the LM relation for our WL selected clusters (red triangles), and compared to the clusters studied in Mahdavi et al. (2013) brown circles.

For the purposes of this comparison the X-ray luminosities are bolometric and in each case the WL masses are determined within a WL defined \( r_{500} \). We fit a power law of the form \( L_X = E(z)^2 L_0(M/M_0)^{B_{LM}} \) to the Mahdavi et al. (2013) clusters finding a normalisation of \( L_0^{\text{LM,13}} = (2.83 \pm 0.47) \times 10^{45} \text{ergs s}^{-1} \) and slope \( B_{LM,13} = 1.89 \pm 0.83 \), with \( M_0 = 8 \times 10^{14} \text{M}_\odot \). We fit to our WL clusters with the slope fixed at the value found for the Mahdavi et al. (2013) relation (\( B_{LM,13} = 1.89 \)), and find a normalisation of \( L_0 = (3.8 \pm 0.8) \times 10^{45} \text{ergs s}^{-1} \) (at \( M_0 = 2 \times 10^{14} \text{M}_\odot \)), see red solid line Fig 6. Therefore, for a given mass we find that our WL selected clusters are \( 5.4 \pm 1.4 \) times under-luminous as compared to the Mahdavi et al. (2013) clusters in terms of the LM relation, with the masses derived from WL and within a WL determined \( r_{500} \) (corrected for centroid position, Eddington bias and triaxiality).

Conversely, the masses are over-estimated by a factor \( 2.4 \pm 0.3 \) for a given luminosity. To validate that the X-ray selected clusters of Mahdavi et al. (2013), whose masses are derived from WL and within \( r_{500} \), follow other X-ray selected clusters, we plot the Pratt et al. (2009) clusters (Fig green squares). Clearly, the two X-ray selected samples are in agreement, even though the mass estimates are from two different methods, with our WL selected clusters offset from both samples. We note that if the Mahdavi et al. (2013) WL masses are calculated within an X-ray derived \( r_{500} \), this has a negligible effect on their LM relation.

Since the discrepancy with the Mahdavi et al. (2013) sample is still prevalent (see Fig 6), even after correcting our WL masses for centroid position, Eddington bias and triaxiality, this implies that there is still some underlying effect due to the WL selection of these clusters.

5.4 Are WL Selected Clusters Under-luminous?

In Section 5.1 we have shown that by considering various systematic effects on the WL cluster mass, we are unable to reconcile the differences seen in the LM relation (see Fig 4), without including corrections for \( r_{500} \) which we argue masks the true discrepancy (see §5.3 Fig 8). Not only that, our previous work (Hamana et al. 2009) found that the WL mass is consistent with previous studies when compared to the velocity dispersion (\( \sigma_v \)). We reproduce this comparison using \( M_{WL,500} \) in Figure 7. We find that the WL clusters (red triangles) are consistent with the X-ray selected clusters studied in Hoekstra (2007) grey crosses, in terms of the \( M_{WL,500} - \sigma_v \) relation. Therefore, we are presented with the situation that the WL mass and \( \sigma_v \) are consistent and

\( ^5 \) The data were downloaded from http://sfstar.sfsu.edu/cccp/
the luminosity and temperature are consistent, implying the ICM properties are inconsistent with \( \sigma_v \).

To quantify this, we derive a \( L_X,\text{bol}-\sigma_v \) relation for the WL clusters. Figure 8 plots the \( L_X,\text{bol}-\sigma_v \) relation for our clusters (red triangles), compared to the relation for the HIFLUGCS sample of clusters (cyan diamonds) studied in Zhang et al. (2011). The cluster SLJ1634 is highlighted by the black circle. Due to the composite nature of this cluster (see §3.3.5), the value of \( \sigma_v \) from Hamana et al. (2009) \( \sigma_v=1402\pm12\ \text{km s}^{-1} \) is likely an overestimate. We therefore re-derive \( \sigma_v \) from our tree analysis using the galaxies from group B (see Fig A9, balck squares), as two of these galaxies are associated with the X-ray emission, finding \( \sigma_v=390\pm220 \text{ km s}^{-1} \). We find that five of the WL selected clusters are clearly under-luminous for their velocity dispersion.

These results imply that there is a strong systematic effect remaining that is influencing the WL and dynamical properties, the ICM properties, or both. We consider the most plausible effect to be a high incidence of projected filamentary structure along the LOS to these WL selected clusters, leading to enhanced WL masses and \( \sigma_v \) relative to halo mass. From the spectroscopic follow-up, many of the clusters show signs of non-gaussianity in their velocity dispersion histograms (see Fig 1 in Hamana et al. 2009), which would be consistent with projected filament structure. This however is based on relatively small numbers of spectra in each case (an average of 14 galaxies per cluster), so a more detailed analysis is unable to be performed.

Bower et al. (1997) investigated a sample of optically selected, X-ray under-luminous clusters and proposed two scenarios for their observed offset in the \( L_X-\sigma_v \) relation. Either the clusters have not yet formed or are at a stage of rapid mass accretion, which would be manifested in a low X-ray luminosity. The second scenario is that the clusters are embedded in filaments viewed along the LOS, which would lead to an over-estimate of \( \sigma_v \) due to contamination of galaxies along the filament. Other previous studies looking at X-ray under-luminous clusters (e.g. Popesso et al. 2007; Balogh et al. 2011; Castellano et al. 2011) favour the scenario that these clusters are still in a stage of formation or that the gas has been expelled. However, since we have measured temperatures for the majority of our clusters and shown that the gas properties scale consistently in the luminosity-temperature relation (and consistent to X-ray selected cluster samples), we favour the scenario that the WL selected clusters are embedded in filaments viewed near the LOS. This is further corroborated by our measurements of the BCG ellipticity, where the majority of our cluster BCGs appear circular, implying that the WL selection strongly favours clusters elongated along the LOS and embedded in correlated filaments.

6 SUMMARY AND CONCLUSIONS

We have presented the results of X-ray observations of 10 WL selected galaxy clusters in order to determine their X-ray properties. Our main results are as follows.

(i) We find that the scaling relation between X-ray luminosity and temperature is consistent with samples of X-ray selected clusters.

(ii) We find that the WL selected clusters are offset from X-ray selected clusters and groups in the luminosity-temperature relation, implying either the X-ray luminosity is under-estimated or the WL mass is over-estimated.

(iii) The centroid shifts for the WL clusters show that the majority appears dynamically unrelaxed compared to X-ray selected samples, suggesting X-ray cluster surveys preferentially detect clusters that are morphologically relaxed and/or WL selection favours morphologically disturbed clusters.

(iv) The cuspliness of the gas density profiles shows that two of the clusters in our sample appear to host a cool core.

(v) Measuring the ellipticities of the BCGs for each cluster, and assuming that they are intrinsically prolate, we find that 60% of the clusters appear to be viewed with their major axis close to the line of sight.

(vi) Through a series of corrections to the WL cluster...
mass including the centroid position, Eddington bias, and triaxiality, and calculating the WL masses with an X-ray derived $r_{200}$, we find the WL selected clusters are consistent in the LM relation with X-ray selected clusters.

(vii) Comparisons to X-ray selected samples with WL derived cluster masses, and within WL defined $r_{200}$, show that the offset in the LM relation is still present, even when correcting for centroid position, Eddington bias, and triaxiality.

Our results show that WL selected clusters are affected by biases which, when combined, has a large effect on the calculated WL mass. We have presented a series of possible biases to explain the large offset of the LM relation when compared to X-ray selected cluster samples. When considering the cumulative effect of these biases on the WL mass we reconcile the difference in the LM relation. Our comparison to an X-ray selected sample of clusters with WL derived cluster masses, show that even after taking into account biases which may arise due to the WL selection, there is still a discrepancy in the LM relation. This implies that there maybe further biases in the WL selection, or that the WL selected clusters are embedded in filamentary structure viewed along the line of sight.

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APPENDIX A: X-RAY AND \textit{SUBARU} IMAGES OF INDIVIDUAL CLUSTERS

\textbf{Figure A1.} Images for the cluster SLJ0225 (z=0.1395). Figure (a) shows the XMM image of the cluster in the 0.7 – 2.0 keV band smoothed by a Gaussian of 1.5 pixel radius (where 1 pixel = 2.2'); (b) shows a \textit{Subaru} image of the cluster within a region \(400\times400''\) in length (shown by the black box in Fig(a)) with X-ray contours (red) and WL mass contours (blue) over-plotted. The X-ray contours were constructed from an adaptively smoothed image of the total MOS data, due to the cluster falling on the chip gaps of the pn-camera. The WL contours are taken from \cite{Hamana et al.} (2009).

\textbf{Figure A2.} Images for the cluster SLJ0850 (z=0.1938). (a) shows a \textit{Chandra} image in the 0.7–2.0 keV band within a region \(25\times25''\) in length, the image is smoothed by a Gaussian of 1.5 pixel radius (where 1 pixel = 1.968''); (b) shows a \textit{Subaru} image of the cluster within a region \(400\times400''\) in length (shown by the black box in Fig (a)) with the X-ray contours (red) and WL mass contours (blue) over-plotted. The X-ray contours were constructed from an adaptively smoothed \textit{Chandra} image of the cluster.
Figure A3. Images for the cluster SLJ1000 (z=0.2166). Figures (a) and (b) same as A2.

Figure A4. Images for the cluster SLJ1135 (z=0.2078). Figures (a) and (b) same as A2 however due to the low SNR of this cluster an adaptively smoothed image of the cluster was not obtained.
Figure A5. Images for the cluster SLJ1204 (z=0.2609). Figures (a) and (b) same as A2.

Figure A6. Images for the cluster SLJ1335 (z=0.4070). Figures (a) and (b) same as A2.
Figure A7. Images for the cluster SLJ1337 (z=0.1798). Figures (a) and (b) same as A2.

Figure A8. Images for the cluster SLJ1602 (z=0.4155). Figures (a) and (b) same as A2.
Figure A9. Images for the cluster SLJ1634 (z=0.2377). Figures (a) and (b) are as in A2. Also highlighted on Figure (b) is the locations of spectroscopically confirmed cluster galaxies, separated into three groups A, B and C (magenta, black and green squares respectively) found by our tree analysis (see §3.3.5).

Figure A10. Images for the cluster SLJ1647 (z=0.2592). Figures (a) and (b) are as in A2.